

REDSHIFT EVOLUTION OF THE DYNAMICAL PROPERTIES OF MASSIVE GALAXIES FROM SDSS-III/BOSS

ALESSANDRA BEIFIORI^{1,2,3}, DANIEL THOMAS^{2,4}, CLAUDIA MARASTON², OLIVER STEELE², KAREN L. MASTERS^{2,4}, JANINE PFORR^{2,5},
ROBERTO P. SAGLIA^{1,3}, RALF BENDER^{1,3}, RITA TOJEIRO², YAN-MEI CHEN^{6,7}, ADAM BOLTON⁸, JOEL R. BROWNSTEIN⁸,
JONAS JOHANSSON^{2,9}, ALEXIE LEAUTHAUD¹⁰, ROBERT C. NICHOL^{2,4}, DONALD P. SCHNEIDER^{11,12}, ROBERT SENGER¹,
RAMIN SKIBBA¹³, DAVID WAKE^{6,14}, KAIKE PAN¹⁵, STEPHANIE SNEDDEN¹⁵, DMITRY BIZYAEV¹⁵, HOWARD BREWINGTON¹⁵,
VIKTOR MALANUSHENKO¹⁵, ELENA MALANUSHENKO¹⁵, DANIEL ORAVETZ¹⁵, AUDREY SIMMONS¹⁵,
ALAINA SHELDEN¹⁵, AND GARRETT EBELKE¹⁵

¹ Max-Planck-Institut für Extraterrestrische Physik, Giessenbachstraße, D-85748 Garching, Germany; beifiiori@mpe.mpg.de

² Institute of Cosmology and Gravitation, University of Portsmouth, Dennis Sciama Building, Burnaby Road, Portsmouth PO1 3FX, UK

³ Universitäts-Sternwarte München, Scheinerstrasse 1, D-81679 München, Germany

⁴ SEPNET, South East Physics Network

⁵ NOAO, 950 N. Cherry Avenue, Tucson, AZ 85719, USA

⁶ Department of Astronomy, University of Wisconsin-Madison, 475 N. Charter Street, Madison, WI 53706, USA

⁷ Department of Astronomy, Nanjing University, Nanjing 210093, China

⁸ Department of Physics and Astronomy, University of Utah, Salt Lake City, UT 84112, USA

⁹ Max-Planck Institut für Astrophysik, Karl-Schwarzschild Straße 1, D-85748 Garching, Germany

¹⁰ Institute for the Physics and Mathematics of the Universe (IPMU), The University of Tokyo, Chiba 277-8582, Japan

¹¹ Department of Astronomy and Astrophysics, The Pennsylvania State University, University Park, PA 16802, USA

¹² Institute for Gravitation and the Cosmos, The Pennsylvania State University, University Park, PA 16802, USA

¹³ Department of Physics, Center for Astrophysics and Space Sciences, University of California, 9500 Gilman Drive, San Diego, CA 92093, USA

¹⁴ Department of Physical Sciences, The Open University, Milton Keynes MK7 6AA, UK

¹⁵ Apache Point Observatory, P.O. Box 59, Sunspot, NM 88349-0059, USA

Received 2013 November 20; accepted 2014 May 6; published 2014 June 18

ABSTRACT

We study the redshift evolution of the dynamical properties of $\sim 180,000$ massive galaxies from SDSS-III/BOSS combined with a local early-type galaxy sample from SDSS-II in the redshift range $0.1 \leq z \leq 0.6$. The typical stellar mass of this sample is $M_{\star} \sim 2 \times 10^{11} M_{\odot}$. We analyze the evolution of the galaxy parameters effective radius, stellar velocity dispersion, and the dynamical to stellar mass ratio with redshift. As the effective radii of BOSS galaxies at these redshifts are not well resolved in the Sloan Digital Sky Survey (SDSS) imaging we calibrate the SDSS size measurements with *Hubble Space Telescope*/COSMOS photometry for a sub-sample of galaxies. We further apply a correction for progenitor bias to build a sample which consists of a coeval, passively evolving population. Systematic errors due to size correction and the calculation of dynamical mass are assessed through Monte Carlo simulations. At fixed stellar or dynamical mass, we find moderate evolution in galaxy size and stellar velocity dispersion, in agreement with previous studies. We show that this results in a decrease of the dynamical to stellar mass ratio with redshift at $>2\sigma$ significance. By combining our sample with high-redshift literature data, we find that this evolution of the dynamical to stellar mass ratio continues beyond $z \sim 0.7$ up to $z > 2$ as $M_{\text{dyn}}/M_{\star} \sim (1+z)^{-0.30 \pm 0.12}$, further strengthening the evidence for an increase of M_{dyn}/M_{\star} with cosmic time. This result is in line with recent predictions from galaxy formation simulations based on minor merger driven mass growth, in which the dark matter fraction within the half-light radius increases with cosmic time.

Key words: galaxies: elliptical and lenticular, cD – galaxies: evolution – galaxies: formation – galaxies: high-redshift – galaxies: kinematics and dynamics

Online-only material: color figures

1. INTRODUCTION

Early-type galaxies play an important role in observational studies of galaxy formation and evolution. Tight empirical correlations between the observed dynamical and stellar population properties of early-type galaxies have been derived that set useful constraints to their formation histories. These are correlations between size (effective radius, R_e), surface brightness and stellar velocity dispersion (σ_e), i.e., the fundamental plane (FP; Dressler et al. 1987; Djorgovski & Davis 1987; Bender et al. 1992, 1993), the stellar mass plane (Hyde & Bernardi 2009b; Auger et al. 2010a), the fundamental manifold of galaxies (Zaritsky et al. 2008), as well as correlations between galaxy color and stellar population age and metal abundance with galaxy mass (see review by Renzini 2006).

Such scaling relations represent a powerful phenomenological tool to study the co-evolution of baryonic and dark matter in galaxies. The latter has been studied extensively for the local galaxy population. The tightness of the FP has been used to constrain stellar population variations or dark matter content in galaxies (Renzini & Ciotti 1993) or to study non-homology (Ciotti et al. 1996). Gerhard et al. (2001) studied the dynamical properties and dark halo scaling relations of giant elliptical galaxies and found that the tilt of the FP is best explained by a stellar population effect and not by an increasing dark matter fraction with luminosity.

Cappellari et al. (2006), using data from the SAURON survey (de Zeeuw et al. 2002), suggested that dynamical to stellar mass-to-light ratios larger than one are due to dark matter, assuming a constant stellar initial mass function (IMF; Kroupa 2001) and

using self-consistent models (see Cappellari et al. 2013a for a review). Other studies based on Sloan Digital Sky Survey (SDSS; York et al. 2000) data came to the conclusion that there is an excess over the predictions of stellar population models with a fixed IMF which are luminosity dependent (Padmanabhan et al. 2004; Hyde & Bernardi 2009a, 2009b).

These conclusions have recently been revised in Cappellari et al. (2013b) by means of a large number of axisymmetric dynamical models including different representations of dark matter halos reporting a variation of IMF slope with galaxy mass. In fact, as highlighted by Thomas et al. (2011), for instance, dark matter fraction and IMF are highly degenerate. They studied a sample of Coma galaxies and their detailed decomposition into luminous and dark matter reveals that for low-mass galaxies there is a good agreement between dynamical masses with dark matter halo and lensing results (galaxies with a $\sigma \sim 200 \text{ km s}^{-1}$ are consistent with a Kroupa IMF). For higher-mass galaxies ($\sigma > 200 \text{ km s}^{-1}$), the disagreement can be due to either a non constant IMF (Kroupa IMF under-predicts luminous dynamical masses for galaxies at $\sigma \sim 300 \text{ km s}^{-1}$) or to a dark matter component which follows the light (see also Wegner et al. 2012).

A promising complementary approach to detailed studies of local galaxies for breaking the degeneracy between dark matter fraction and IMF is to study the evolution of fundamental plane, dynamical and stellar population properties of galaxies with look-back time (Bender et al. 1998; van Dokkum et al. 1998; Treu et al. 2005; Jørgensen et al. 2006; Saglia et al. 2010; Houghton et al. 2012; Bezanson et al. 2013a). Moreover, dark matter fractions can also be studied in samples of lensed galaxies at different redshifts (see Bolton et al. 2012a using data from both the Sloan Lens ACS sample, SLACS; Bolton et al. 2006, 2008, and the BOSS Emission line Lens Survey, BELLS; Brownstein et al. 2012).

A large number of such investigations have been performed in recent years, analyzing the redshift evolution of galaxy sizes (e.g., Daddi et al. 2005; Trujillo et al. 2006a, 2006b, 2007; Longhetti et al. 2007; Zirm et al. 2007; Toft et al. 2007; van Dokkum et al. 2008; Buitrago et al. 2008; Cimatti et al. 2008; Franx et al. 2008; Bernardi 2009; Damjanov et al. 2009; Saracco et al. 2009; Bezanson et al. 2009; Mancini et al. 2009, 2010; Valentinuzzi et al. 2010a; Carrasco et al. 2010; Szomoru et al. 2012; Newman et al. 2012; Saracco et al. 2014) and dynamical properties of galaxies (van der Wel et al. 2005, 2008; Cenarro & Trujillo 2009; Cappellari et al. 2009; van Dokkum et al. 2009; Newman et al. 2010; Onodera et al. 2010; Saglia et al. 2010; van de Sande et al. 2011, 2013; Toft et al. 2012; Onodera et al. 2012; Damjanov et al. 2013; Belli et al. 2014). The bottom line is that galaxy sizes appear to decrease (at fixed stellar mass) and stellar velocity dispersions increase (at fixed stellar mass) with increasing look-back time. Some of these conclusions are still controversial, however. Tiret et al. (2011), for example, argue that the size evolution disappears when one homogenizes literature data sets by measuring the stellar mass with the same method and provides accurate estimates of the sizes to prevent systematic effects in the R_e measurements of low signal-to-noise ratio (S/N) high- z compact galaxies (Mancini et al. 2009). Mancini et al. (2010) also report evidence for galaxies as large as local ones at redshift higher than 1.4 suggesting that not all high- z galaxies are compact. Moreover, the effect depends on the photometric band and on whether galaxies have young or old light-weighted ages and if they reside in clusters or field (see Valentinuzzi et al. 2010b for clusters and Poggianti et al.

2013 for field studies; see also Trujillo et al. 2011 for a different opinion).

So far, galaxies in the distant universe (look-back times of a few billion years and above) have not been studied at the same statistical level as local galaxies. The new data set from the Baryon Oscillation Spectroscopic Survey (BOSS; Dawson et al. 2013) as part of the Sloan Digital Sky Survey-III (SDSS-III; Eisenstein et al. 2011) provides the opportunity to investigate the dynamical and stellar population properties of a galaxy sample of unprecedented size up to redshifts $z \sim 0.7$. The survey is currently obtaining spectroscopic data for nearly 1.5 million massive galaxies at redshifts between 0.2 and 0.7, hence up to look-back times of ~ 6 Gyr. The combination of this sample with local early-type galaxies from SDSS-I/II allows us to make a statistically significant link between local galaxy properties and higher redshift observations. This is the main aim of the present paper.

The paper is organized as follows. The galaxy sample is described in Section 2. Photometric, kinematic, and dynamical properties are presented in Section 3, as well as the calibration technique for measuring the effective radii based on a sub-sample of galaxies with *Hubble Space Telescope* (HST) photometry and the correction for progenitor bias. The results are presented in Section 4 and discussed in Section 5. The paper concludes with Section 6.

Throughout the paper, we assume the following cosmology with $H_0 = 71.9 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $\Omega_m = 0.258$, and $\Omega_\Lambda = 0.742$ following the cosmology used for the stellar mass determination of BOSS galaxies in DR9 (Maraston et al. 2013).

2. DATA

We use the galaxy sample from SDSS-III/BOSS covering a redshift range $0.2 \lesssim z \lesssim 0.7$. To leverage our study on the redshift evolution, we combine this sample with a local sample of massive galaxies at $z \sim 0.1$ drawn from SDSS-II.

2.1. Main Galaxy Sample from SDSS-III/BOSS

Data are taken from the SDSS-III/BOSS Data Release Nine (DR9, Ahn et al. 2012). BOSS (Dawson et al. 2013; Smee et al. 2013), one of the four surveys of SDSS-III (Eisenstein et al. 2011), is taking spectra for 1.5 million luminous massive galaxies with the aim of measuring the cosmic distance scale and the expansion rate of the Universe using the Baryonic Acoustic Oscillations (BAO) scale (Anderson et al. 2012). Data are taken with an upgraded version of the multi-object spectrograph on the SDSS telescope (Gunn et al. 2006). BOSS galaxy targets are selected from the SDSS *ugriz* imaging (Fukugita et al. 1996; Gunn et al. 1998; Stoughton et al. 2002), including new imaging part of DR8 mapping the southern Galactic hemisphere. A series of color cuts have been used to select targets for BOSS spectroscopy (Dawson et al. 2013). The selection criteria are designed to identify a sample of luminous and massive galaxies with an approximately uniform distribution of stellar masses following the Luminous Red Galaxy (LRG; Eisenstein et al. 2001) models of Maraston et al. (2009). The galaxy sample is composed of two populations: the higher-redshift Constant Mass Sample (CMASS; $0.4 < z < 0.7$) and the Low-Redshift Sample (LOWZ; $0.2 < z < 0.4$). A fraction (around one-third) of those LOWZ galaxies derived with those cuts have been already observed in SDSS-I/II and are included in the BOSS sample, but they are not re-observed if they had reliable

redshifts. These two population are well separated in the $(g - r)$ and $(r - i)$ colors diagram (Masters et al. 2011).

To extract a working set of objects from the entire BOSS sample, we matched galaxies from different catalogs of stellar velocity dispersions and stellar masses described in the following sections using the keywords PLATE, MJD, FIBERID that uniquely determine a single observation of a single object. The final merged catalog comprises 491,954 galaxies that are included in DR9. We selected objects from the LOWZ and CMASS samples (BOSS_TARGET1 = Target flags) with a good platequality (PLATEQUALITY = ‘good’), with object class ‘galaxy’ (class_noqso = GALAXY), with high-confidence redshifts (ZWARNING_NQSO = 0), and for which we have a unique set of objects in the case of duplicate observations (SPECPRIMARY = 1).

The final sample we will analyze comprises $\sim 180,000$ objects (37% of the original sample) obtained after applying some additional redshift cuts and quality cuts to stellar velocity dispersions and stellar masses that we will discuss in Section 3.

2.2. Local Galaxy Sample from SDSS-II

In order to connect our BOSS galaxies to the local galaxy population, we use a sample of galaxies from the SDSS Data Release Seventh (DR7; Abazajian et al. 2009). We select early-type galaxies following Hyde & Bernardi (2009a). Galaxies had to be well fitted by a de Vaucouleurs profile in the g and r bands ($\text{fracDev}_g = 1$ $\text{fracDev}_r = 1$), with an early-type-like spectrum ($\text{eClass} < 0$), extinction-corrected r -band de Vaucouleurs magnitudes in the range $14.5 < \text{devMag}_r < 17.5$ (this results in a tighter limit compared to that of the SDSS Main Galaxy sample; see details in Hyde & Bernardi 2009a), measured stellar velocity dispersion in DR7 $\text{velDisp} > 0$, and with an axis ratio in the r band of $b/a > 0.6$, to be more likely pressure supported. As described in Hyde & Bernardi (2009a), the DR7 b/a distribution shows two distinct populations separated by this axis-ratio value with a 20% of low-luminosity objects at $b/a < 0.6$.¹⁶ This retains $\sim 123,500$ DR7 objects.

Bernardi et al. (2010) did a detailed comparison of different methods to select early types in the literature (morphologically based, colors, or structural parameters-based methods), showing that Hyde & Bernardi (2009a) cuts give similar results to other methods but is more efficient in discriminating elliptical galaxies from spirals, which is important in our analysis.

We select galaxies with $\text{zWarning} = 0$ and apply some constraints on the errors of the parameters: errors on $R_e < 70\%$, errors on the axis ratios $0 < \text{err}_{b/a} < 1$, errors in $\sigma < 30\%$. Only galaxies with stellar velocity dispersion of $70 < \sigma < 550 \text{ km s}^{-1}$ were selected, following the BOSS cuts.

3. GALAXY PROPERTIES

The primary aim of this work is to study the redshift evolution of the dynamical properties of BOSS galaxies. Stellar masses are taken from Maraston et al. (2013) and stellar velocity dispersions from Thomas et al. (2013). Both of these quantities are included

¹⁶ We do not apply this cut in our BOSS sample because b/a from the SDSS imaging could be highly unreliable. In fact, we do not find the same clear separation in the b/a distribution of BOSS galaxies. Also, the BOSS b/a distribution looks different probably due to the large uncertainties on b/a from SDSS imaging. However, we assess the typical percentage of galaxies which should have a $b/a < 0.6$ by using the sub-sample of BOSS galaxies with COSMOS photometry (see Section 3.2.2 for details) and find that $\sim 22\%$ of galaxies have a $b/a < 0.6$, which is consistent with Masters et al. (2011) findings. Finally, the choice of this cut does not affect appreciably our results.

in the DR9 data release. We used a sub-sample of 240 galaxies with additional *HST*/COSMOS photometry and BOSS spectra (Masters et al. 2011) to derive a calibration for galaxy sizes from DR8. In the following sections, we describe the galaxy parameters in more detail. In Section 3.4, we describe how the effective radii and stellar velocity dispersion are combined to derive virial masses.

The redshifts used in our analysis are those extracted from DR9 as z_noqso with formal 1σ error given by z_err_noqso (outputs of the BOSS pipeline as described in Bolton et al. 2012b; Dawson et al. 2013). Redshifts are successfully determined for $\sim 98\%$ of CMASS galaxies (Anderson et al. 2012, their Table 1). Errors in the measured redshift are less than about 0.0002 ($\sim 60 \text{ km s}^{-1}$). BOSS observed a few galaxies at $z > 0.7$ and $z < 0.2$. In our analysis, we focus on the redshift range $0.2 \leq z \leq 0.7$ and therefore excluded all the galaxies outside this redshift range. This cut retains 90% of the galaxies (444,118 objects).

3.1. Stellar Mass

We use the stellar masses M_* from Maraston et al. (2013) as published in the DR9 data release.¹⁷ These masses are derived through broadband spectral energy distribution (SED) fitting of model stellar populations on SDSS u, g, r, i, z modelMag magnitudes from DR8, scaled to the i -band cmode1Mag magnitude. The BOSS spectroscopic redshift is used to constrain the fits. We note that the DR9 data release also provides alternative mass estimates from Chen et al. (2012; see Maraston et al. 2013 for discussion).

Maraston et al. (2013) use different types of templates to derive stellar masses; passive models, star-forming models or a mix, chosen so as to match the galaxy expected galaxy type, based on a color cut, in BOSS.¹⁸ In this paper we used stellar masses obtained with the passive template that is a mix of old populations with a spread in metallicity and which was found to reproduce well the colors of luminous red galaxies at redshift 0.4 to 0.7 (Maraston et al. 2009). This maximally old and passive LRG template minimizes the risk of underestimating M_* which is the typical effect that occurs when one determines stellar masses from light (Maraston et al. 2010; Pforr et al. 2012).

We adopt the stellar masses derived from the median of the probability distribution function (PDF). Typical uncertainties on M_* are < 0.1 dex, independently of redshift (see Maraston et al. 2013 for details).

The Portsmouth stellar mass pipeline described in Maraston et al. (2013) provides stellar masses using various types of configurations, i.e., Kroupa (2001) or Salpeter (1955) IMF, passive (Maraston et al. 2009) or star-forming (Maraston et al. 2006) models, and considering or not the mass loss in the stellar evolution. The subset of calculations used here adopt a Kroupa IMF. As is well known, a Salpeter IMF produces systematically larger stellar masses by about a factor 1.6 (Maraston 2005; Bolzonella et al. 2010). Systematic uncertainties are mainly due to the choice of the IMF (~ 0.2 dex).

We removed galaxies for which M_* was not properly determined due to unreliable values of redshift and photometry issues. This cut retains 387,590 galaxies (79% of the full sample).

¹⁷ Stellar masses have been derived for galaxies with non-zero photometry in $i\text{-band mode1mag}_i > 0.0$, $\text{z_err_noqso} \geq 0$ and $\text{z_noqso} > \text{z_err_noqso}$.

¹⁸ This assignment of passive or star-forming models following a color cut is strictly valid at $z > 0.4$. However, since LOWZ galaxies are generally red, Maraston et al. (2013) assumed the criterion valid over the full BOSS range.

3.2. Size

The photometric data used in this paper were derived using the SDSS-III DR8 pipeline (Aihara et al. 2011). One of the main updates performed in DR8 compared to Stoughton et al. (2002) is the correction for sky levels (in particular for extended galaxies) which past studies found to be overestimated (Bernardi et al. 2007; Lauer et al. 2007; Guo et al. 2009). This issue has also been addressed in detail in Blanton et al. (2011). SDSS effective radii are adopted from the DR8 catalog. In this section we describe this measurement and its calibration with COSMOS/*HST* imaging.

3.2.1. SDSS DR8 Effective Radii

Effective radii are estimated using seeing-corrected de Vaucouleurs effective radii (deVRad) and the associated errors (deVRadErr; of the order of $\sim 5\%$ – 25% depending on redshift). Those values correspond to the effective radius along the semi-major axis derived on elliptical aperture.

BOSS galaxies are often not well resolved in SDSS imaging. The average seeing of the SDSS survey is $1''.05$ from the PSF_FWHM in the *i*-band (lower and upper percentile of $0''.88$ to $1''.24$), which is better than in previous data releases due to the repeated images taken during the survey (see Ross et al. 2011; Masters et al. 2011). The median effective radius of BOSS galaxies (after having applied the previous cuts on the sample) is $1''.24$ (upper and lower percentile $2''.12$ and $0''.72$, respectively). As a consequence, seeing effects will affect size measurements (see Saglia et al. 1993, 1997; Bernardi et al. 2003), which we need to correct for. To develop a seeing correction we compare SDSS galaxy radii with measurements based on high-resolution *HST*/COSMOS imaging (Section 3.2.3).

The surface brightness models used by pipeline to obtain galaxy sizes are relatively simple (single parameter fits, i.e., exponential or de Vaucouleurs profiles); a more complex model is not feasible for this analysis because of the limitation of the image resolution. We tested this conclusion and found that performing Sérsic fits on BOSS images yields strong degeneracies between effective radius and Sérsic index, preventing robust estimate of effective radii. We estimate SDSS circularized radii as $R_{e,circ} = R_e \times q^{1/2}$ where R_e is the semi-major axis of the half-light ellipse and $q = b/a$ is the axis ratio which is part of the DR8 catalog.

The effective radii should, in principle, be referred to a fixed rest-frame wavelength to account for the fact that early types have color gradients, so on average their optical radii are larger in bluer bands and at higher z this effect will make the sizes larger (see Bernardi et al. 2003; Hyde & Bernardi 2009a for discussions). We did not apply any correction for this trend, because uncertainties of the size calibration ($\sim 10\%$ – 25%) are larger than the typical variation in size measured in different filters (from 4% to 10%; Bernardi et al. 2003; Hyde & Bernardi 2009a). In Section 4.1 we describe the average effect this could have on the size evolution studies.

Effective radii were converted to physical radii using the code of Hogg (1999) which presents the scale conversion between arcsec and kiloparsec for our given cosmology. Hereafter, we will use the notation *pipeline* R_e to represent SDSS effective radii in kiloparsecs. We further remove from the catalogs galaxies with unreliable values of R_e and their errors. This cut retains $\sim 370,000$ objects.

3.2.2. COSMOS Effective Radii

Masters et al. (2011) constructed a sub-sample of 240 BOSS galaxies with *HST*/COSMOS imaging (Koekemoer et al. 2007; Leauthaud et al. 2007). This sample was used to calibrate the SDSS DR8 radii part of the CMASS sample. We adopted the effective radii from the public Zurich Structure & Morphology Catalog v1.0¹⁹ (Scarlatà et al. 2007; Sargent et al. 2007) derived from *I*-band (F814W) ACS images. Effective radii are available from this catalog for 224 of the 240 objects.

The catalog contains galaxy structural parameters derived from a two-dimensional decomposition using the GIM2D code (Galaxy Image 2D; Simard 1998) on the ACS/*HST* images as deep as $I_{AB} \sim 22.5$ (Sargent et al. 2007). We used the seeing-corrected effective radii R_{GIM2D} resulted from the one component Sérsic fit, of the form

$$I(r) = I_e 10^{[-b_n((r/R_{e,n})^{1/n} - 1)]}, \quad (1)$$

where I_e is the effective intensity, and the constant b_n is defined in terms of the shape parameter n and is chosen so that $R_{e,n}$ encloses half of the total luminosity and it is measured in arcsec. The quantity b_n can be well approximated by $b_n = 0.868n - 0.142$ (Caon et al. 1993). The Zurich Structure & Morphology Catalog v1.0 contains two values of the effective radius from GIM2D, R_{GIM2D} and R_{OP5_GIM2D} which correspond to the point-spread function (PSF)-corrected effective radii. The choice of one of the two is purely arbitrary and they show a negligible median difference of $0''.001$ (with R_{GIM2D} being smaller than R_{OP5_GIM2D}). The catalog gives the uncertainties on effective radii as 99% confidence lower and upper error on R_{GIM2D} (LE_ R_{GIM2D} , UE_ R_{GIM2D} , respectively). Those errors are $< 1\%$ of R_{GIM2D} , much smaller than the errors in R_e from SDSS photometry ($\sim 15\%$ in the COSMOS/BOSS sub-sample).

Sizes are converted to circularized effective radii $R_{e,serc}$ following the procedure described in Saglia et al. (2010)²⁰ which calculates the half-light radius obtained from the classical curve of growth analysis of the intrinsic Sérsic profile. The procedure requires R_{GIM2D} , Sérsic index $SERSIC_N_GIM2D$, and axis ratio $q = b/a$, where a , b are the semi-major and minor axis of the half-light ellipse, and $q = 1 - \epsilon$, where ϵ is the ellipticity of the object ELL_GIM2D . Here we use a different parameterization than in Section 3.2.1 because we have additional information from the Sérsic fit (the two parameterization would give consistent results for a wide range of q).

We did not apply any correction to account for the fact that rest-frame wavelengths are different at different redshifts using only *i*-band data because this sample is only used for calibration purposes, and uncertainties due to the size calibration we derive using COSMOS radii are much larger than the wavelength variation at the redshifts considered (see Section 3.2.1). Moreover, COSMOS has just one filter available for those galaxies. COSMOS radii were converted to kiloparsecs in the same way as SDSS radii using the code of Hogg (1999). Hereafter, we will use the notation *COSMOS* $R_{e,ser}$ to represent COSMOS effective radii in kiloparsecs.

¹⁹ Available at <http://irsa.ipac.caltech.edu/data/COSMOS/datasets.html>.

²⁰ The code is available at http://www.mpe.mpg.de/~saglia/rps_software.html.

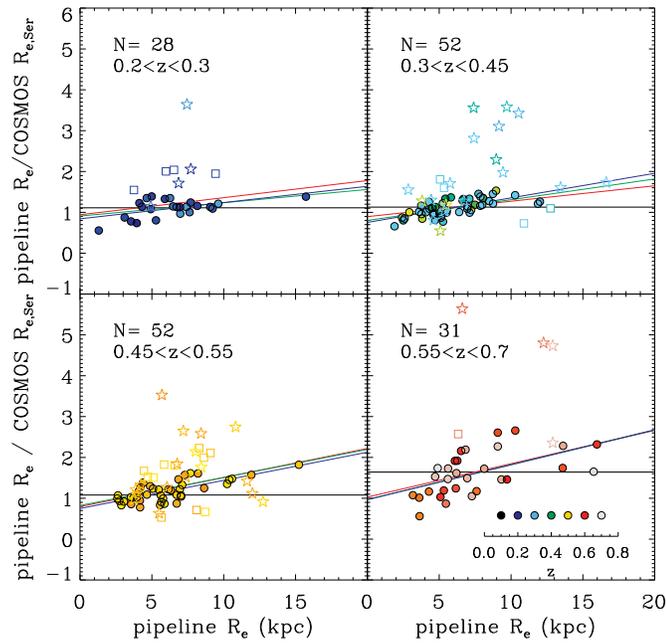


Figure 1. Ratio between pipeline R_e from SDSS DR8 and COSMOS $R_{e,ser}$ as a function of pipeline R_e . Points are coded as a function of redshift ($0.2 \leq z \leq 0.3$, $0.3 < z \leq 0.45$, $0.45 < z < 0.55$, and $0.55 \leq z \leq 0.7$). Different lines represent different fitting procedures: the red line is a linear fit, the green line is a linear fit with one 2σ clip, the blue line is iterative 2σ clipping, the horizontal black line is the single offset. Circles are the points that have been used in the fit and open squares are points discarded in the iterative sigma clipping. Open stars are multiple systems in *HST* imaging which are unresolved in the SDSS images (Masters et al. 2011). Labels in each panel give the number of galaxies used in the fit including objects discarded in the iterative 2σ clipping (multiple systems are not included in this number).

(A color version of this figure is available in the online journal.)

3.2.3. Comparison SDSS versus COSMOS

A significant fraction (23%) of BOSS galaxies with *HST* imaging are unresolved multiple systems in SDSS imaging (Masters et al. 2011). To derive an accurate calibration, we excluded these unresolved multiple systems from our analysis (44 galaxies) using the public catalog of Masters et al. (2011).²¹ The redshift range we wish to study is $0.2 \leq z \leq 0.7$, and we discarded 4 additional galaxies at higher redshifts and 13 at $z < 0.2$. This leaves us a sample of 163 galaxies. COSMOS redshifts are adopted, as not all of these objects (158 out of 163) have BOSS redshifts (this will not change our results, because the median difference between COSMOS and BOSS redshifts is negligible, 7.21×10^{-6}). We do not correct for PSF variations in the SDSS imaging, because the PSF is reasonably stable and the effect is negligible compared to the overall correction derived here.

Figure 1 shows the ratio between SDSS and COSMOS effective radii as a function of SDSS effective radius for four redshift bins. The ratio is around 1.1, and SDSS radii can be overestimated by up to a factor two. The discrepancy between SDSS and COSMOS radii increases with increasing SDSS radius (see also Masters et al. 2011). Masters et al. (2011)²² found that a single offset was reproducing their data in which they compared the ratio SDSS over COSMOS radii as a function of COSMOS radii. In this work, we compare the ratio of SDSS over COSMOS radii as a function of SDSS radii to derive a

²¹ Available at <http://www.icg.port.ac.uk/~mastersk/BOSSmorphologies/>.

²² Masters et al. (2011) estimated the size correction only for the CMASS sample and used major axis radii.

Table 1
Size Correction for the Four Redshift Bins

z Range	N	a	b	c
$0.2 \leq z \leq 0.3$	24	0.84 ± 0.11	0.04 ± 0.01	0.29 ± 0.04
$0.3 < z \leq 0.45$	48	0.75 ± 0.06	0.06 ± 0.01	0.19 ± 0.02
$0.45 < z < 0.55$	41	0.75 ± 0.04	0.07 ± 0.01	0.32 ± 0.03
$0.55 \leq z \leq 0.7$	30	0.97 ± 0.17	0.08 ± 0.02	0.49 ± 0.06

Notes. A correlation of the form $R_{e,pipeline}/R_{e,COSMOS,ser} = a + b(R_{e,pipeline})$ is assumed. N is the number of points used in the fit after iterative 2σ clipping. Uncertainties on each parameter are 1σ errors. The rms scatter c is derived as deviation of the data about the fits considering also objects discarded by the 2σ clipping.

correction for the full BOSS sample. As is to be expected, there is also some redshift dependence, in the sense that SDSS radii overestimate most the true radii at higher redshifts. Also, as expected, the scatter of the relationships increases with redshift owing to the decrease in SDSS imaging quality.

The size calibration could be affected by the larger uncertainties in the SDSS effective radii due to the higher than typical sky background ($\geq 60\%$ – 70%) of SDSS images in the COSMOS field (Masters et al. 2011; Mandelbaum et al. 2012; on the other hand, seeing is smaller than typical of 10% – 15%) which could give relations between SDSS and COSMOS radii not universal for the full BOSS sample.

We performed fits to these relationships in the four redshift bins independently. We tested for linear correlations applying different levels of sigma clipping in the linear fits: no sigma clipping, red line in Figure 1; just one 2σ clipping, green line in Figure 1; and an iterative 2σ clipping, blue line in Figure 1 with a maximum of three iterations.

We fitted a linear relation of the form $R_{e,pipeline}/R_{e,COSMOS,ser} = a + b(R_{e,pipeline})$. The best-fit quantities a and b , the number of galaxies used in the fit after the sigma clipping, the scatter of the relations c (which include objects discarded by the 2σ -clipping), and their associated errors obtained as 1σ uncertainties for each redshift bin are listed in Table 1. The least-square fits were performed using the MPFIT algorithm (Markwardt 2009) under the IDL²³ environment. Fits with and without sigma clipping are consistent within the errors.

The slope of the relation increases slightly with redshift as to be expected. The scatter about the relation is comparable in the first three redshift bins, while the last redshift bin shows a considerably higher scatter (see Table 1; $c = 0.49 \pm 0.06$). For this reason, we will only use the first three redshift bins in the analysis.

We tested the significance of the fits through an F -test by comparing the resulting χ^2 values for free and fixed slope fits accounting for the number of degrees of freedom, and find a maximum probability of no relation to be $\sim 2\%$, which confirms the statistical validity of our fits. The final fits we adopt for the radius correction in each redshift bin are the linear fits with iterative 2σ clipping (blue lines) because they give corrections with a smaller scatter compared to other fits (of 6% – 30%). The open squares in Figure 1 are those galaxies that were discarded in the sigma clipping. Open stars represent unresolved multiple systems not considered in the fits.

We additionally searched for correlations of the effective radius with several other DR8 structural parameters like axis ratio b/a , $fracdev$, and the difference between $fiber2mag$

²³ Interactive Data Language is distributed by Exelis Visual Information Solutions. It is available from <http://www.exelisvis.com/ProductsServices/IDL.aspx>.

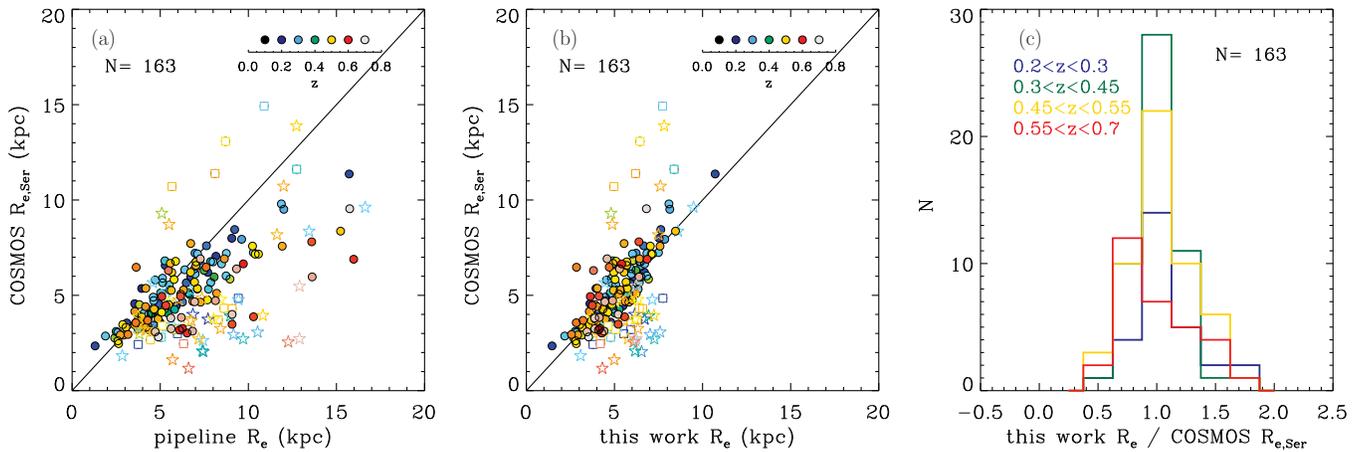


Figure 2. Left panel: COSMOS $R_{e,ser}$ as a function of pipeline SDSS R_e coded in terms of redshift. Middle panel: COSMOS $R_{e,ser}$ as a function of rescaled SDSS R_e (used in the present work) coded in terms of redshift. Symbols are as in Figure 1. Right panel: distribution of the ratio between SDSS rescaled R_e and COSMOS $R_{e,ser}$ for the four redshift bins in Figure 1. Histograms contain also discarded objects but not multiple systems. Legends in each panel give the number of galaxies used in the derivation of the size correction including objects discarded in the iterative 2σ clipping (multiple systems are not included in this number).

(A color version of this figure is available in the online journal.)

and `modelmag` with the aim at finding the best parameter space for the radius correction. None of these parameters helped improving the radius correction.

The size correction we derive here accounts also for the fact that galaxies in our sample could have been better described by a Sérsic profile rather than a de Vaucouleurs, therefore we should consider our calibrated sizes as “Sérsic-like.”

3.2.4. Radius Calibration

Figure 2, panel (b), shows the final corrected radii that are obtained using the fits of Figure 1. For comparison, panel (a) shows the uncorrected radii. Panel (c) presents the distribution of the ratio between the SDSS and COSMOS radii in various redshift bins after the correction. The radii agree well at all redshifts after the correction has been applied. More specifically, the median ratio between our rescaled R_e and COSMOS $R_{e,ser}$ is 1.02 (upper and lower quartile 1.34 and 0.86 and mean 1.10) for $0.2 \leq z \leq 0.3$, 0.98 (upper and lower quartile 1.15 and 0.85 and mean 1.01) for $0.3 < z \leq 0.45$, 1.01 (upper and lower quartile 1.31 and 0.84 and mean 1.05) for $0.45 < z < 0.55$, and 0.99 (upper and lower quartile 1.41 and 0.73 and mean 1.02) for $0.55 \leq z \leq 0.7$. Median values of the distributions in each redshift are compatible within $\pm 1\sigma/N^{1/2}$, where N is the number of objects. Typical errors on rescaled radii range from 0.7 kpc at $z \sim 0.2$ to 1.0 kpc at $z \sim 0.7$ and median radii range from 5.34 to 4.28 kpc, mode 4.73 to 3.48 kpc, in this redshift range. If we did not apply the size correction, we would have larger radii (median sizes range from 5.72 kpc at $z \sim 0.25$ to 5.40 kpc at $z \sim 0.55$, 6.51 kpc at $z \sim 0.65$, mode from 5.00 to 4.00 kpc and 5.00 at $z \sim 0.65$).

Filled circles in Figure 2 are those galaxies that were used in the previous section to derive the calibration. Open squares are those objects that were discarded in the iterative 2σ clipping. Open stars are the unresolved multiple systems discarded from the calibration. Most of the multiples are strong outliers in these plots and they would have been discarded during the sigma clipping fit.

Figure 3 shows the resulting distributions of galaxy effective radii (in both arcseconds and kiloparsecs) for the final sample of 369,254 galaxies in the various redshift bins. The size distribution can be described by a log-normal function (as

the typical size distribution at low redshift; Shen et al. 2003; Bernardi et al. 2003) but with different peaks of the distributions suggesting a variation of typical sizes with z . In Section 4, we present our results using both the corrected SDSS size and pipeline sizes (which we circularized using SDSS axis ratio for this purpose).

3.2.5. Systematic Errors

The systematics in the error budget have been assessed through Monte Carlo simulations which account for uncertainties in both parameters a and b of the fit. The average errors vary with redshift in a non-linear fashion. The errors are ~ 0.6 kpc at $z \sim 0.25$, ~ 0.2 kpc at $z \sim 0.55$, and ~ 0.6 kpc at $z \sim 0.65$.

We additionally include in our Monte Carlo simulations the impact of unresolved multiple systems, which have systematically overestimated sizes. By using the COSMOS/BOSS subsample, we can estimate that those correspond to the $\sim 6\%$ of the galaxies in this sample (see left panel of Figure 13). The sizes of the two components which are resolved in the COSMOS imaging (and unresolved in the SDSS imaging) allow us to assess the contribution of unresolved multiple systems in our analysis, which seems to be negligible compared to other systematic uncertainties (see Appendix B for details).

More detail on the Monte Carlo simulations are given also in Section 4, where we discuss the impact on the final science analysis.

3.3. Stellar Velocity Dispersion

Stellar velocity dispersions (σ) are taken from the Portsmouth Spectroscopic pipeline described in Thomas et al. (2013), also available in DR9. Briefly, stellar kinematics are derived by means of the Penalized Pixel-Fitting method pPXF (Cappellari & Emsellem 2004) in spectra in which emission lines are fitted with Gaussian templates by using the GANDALF code (Sarzi et al. 2006). The stellar population models of Maraston & Strömbäck (2011) have been adopted to fit the stellar continuum. These are based on a hybrid model between MILES stellar library (Sánchez-Blázquez et al. 2006) and theoretical spectra at bluer wavelengths from UVBLUE (Rodríguez-Merino et al. 2005). Stellar population models based on the MILES library have a resolution of 2.54 \AA FWHM (Beifiori et al. 2011), and

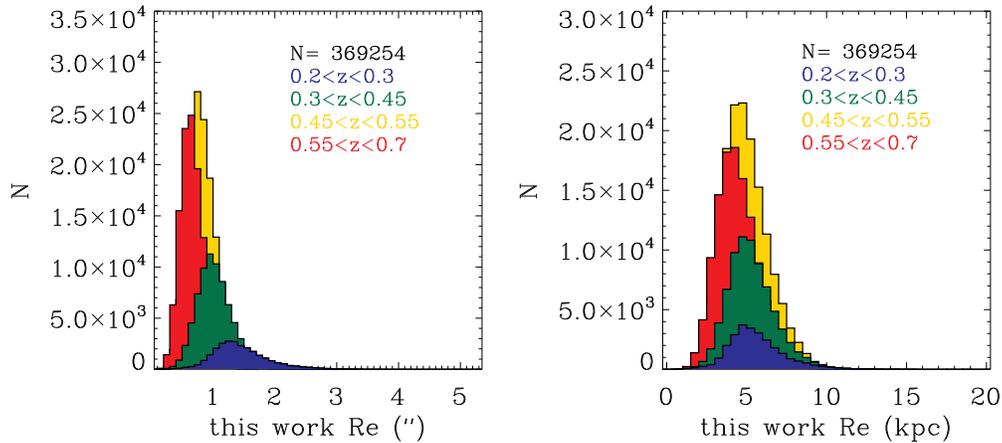


Figure 3. Histogram of the effective radii derived in this work (after size correction), in arcseconds (left panel) and kiloparsecs (right panel) for various redshift bins as color-coded in Figure 2.

(A color version of this figure is available in the online journal.)

therefore needed to be only slightly downgraded to match the BOSS spectral resolution ($R \sim 1800\text{--}2000$ at 5000 \AA , 2.78 \AA -2.50 \AA FWHM). Stellar velocity dispersions have been measured in the typical rest-frame wavelength range $4500\text{--}6500 \text{ \AA}$ most suitable for stellar kinematics analysis due to the presence of strong absorption features (Bender 1990; Bender et al. 1994).

Stellar velocity dispersions from the Portsmouth Spectroscopic pipeline agree within a few percent with other DR9 measurements of σ by Bolton et al. (2012b) and Chen et al. (2012; see Thomas et al. 2013 for a detailed comparison of the systematic offsets between methods). Thomas et al. (2013) show that the typical error distribution on the σ measurements for BOSS galaxies peaks at 14%, and 93% of the measurements have an error below 30%. We therefore selected objects with an error in σ below 30% for the present study to be as inclusive as possible while still maintaining an acceptable accuracy in velocity dispersion (large errors are due to the low S/N of BOSS spectra, mean ~ 4.4 from S_N median, which is sufficient to measure velocity dispersions; Thomas et al. 2013). This cut is not as tight as is generally applied but it allows us not to be affected by biases due to sample selection (for example, a common tighter cut with a relative error $< 10\%$ would discard most of the low σ galaxies at high redshift). Thomas et al. (2013) also show that σ determinations show no bias with S/N. Errors on σ_e slightly vary with redshift, from 12 km s^{-1} at $z \sim 0.25$ to 39 km s^{-1} at $z \sim 0.65$.

Besides the cut in relative error below 30%, we further restrict our sample to values of $70 \leq \sigma \leq 550 \text{ km s}^{-1}$. We discard velocity dispersions below 70 km s^{-1} because of the limit in instrumental resolution of the BOSS spectrograph, and velocity dispersions above 550 km s^{-1} to exclude contamination by potential multiple systems (Bernardi et al. 2003, 2006, 2008). The final number of galaxies that survive these additional cuts in velocity dispersion is $\sim 370,000$, which is 75% of the original sample.

The stellar velocity dispersions from BOSS spectroscopy (σ_{ap}) are measurements within the $2''$ diameter aperture of the BOSS fiber. Therefore, we applied an aperture correction to translate the BOSS velocity dispersions to the aperture corresponding to the effective radius using the relation of Cappellari et al. (2006) derived from the integral field data of the SAURON sample

$$\sigma_e = \sigma_{\text{ap}} \times (r_{\text{ap}}/R_{e,*})^{0.066} \quad (2)$$

in which σ_e is the stellar velocity dispersion within $R_{e,*}$, and $r_{\text{ap}} = 1''$ is the radius of the BOSS fiber. $R_{e,*}$ is taken from the rescaled effective radii converted to arcsecond. The relation of Cappellari et al. (2006) is consistent with that of Mehlert et al. (2003, slope = 0.06) and slightly steeper but in agreement within the errors with older determinations by Jorgensen et al. (1995, slope = 0.04).

Aperture corrections depend on galaxy profile and systematic evolution in the light profile of galaxies could affect the stellar velocity dispersion, as well as this rescaling factor could change from local SAURON galaxies to the higher redshift BOSS galaxies. However, we expect this effect to be negligible as the aperture corrections are small (maximum 3% at higher redshift) because the fiber diameter is close to the typical effective radius of galaxies at the redshifts studied here (see Figure 3). Typical uncertainties after the aperture correction range from 5% to 16% of σ_e (13 to 39 km s^{-1}).

3.3.1. Systematic Errors

We performed Monte Carlo simulations to estimate systematic errors on the aperture correction due to the size calibration (see Section 3.2.4), and have found them to be small. On average σ_e changes by $\sim 1.5 \text{ km s}^{-1}$, which is well below the measurements errors.

3.4. Dynamical Mass

Following Beifiori et al. (2012), we estimate dynamical galaxy mass from the effective radius and velocity dispersion within the effective radius using the virial mass estimator as

$$M_{\text{dyn}} = \beta_{\text{dyn}}(n) R_{e,*} \sigma_e^2 / G, \quad (3)$$

where G is the gravitational constant and β_{dyn} is a dimensionless constant that depends on galaxy structure, often adopted as $\beta_{\text{dyn}} = 5.0 \pm 0.1$ for local galaxies; see Cappellari et al. (2006).

Even though based on measurements within the effective radius, this virial estimator is designed to capture the total dynamical mass of a galaxy. A caveat is that this might only be true as long as total mass traces light. Thomas et al. (2011) found that this assumption might not be consistent with lensing studies. They suggest that Equation (3) may only yield about 86% of the true total dynamical mass. However, this will only affect the absolute scale of the dynamical to stellar mass ratios that we derive, while their evolution with redshift will remain the same.

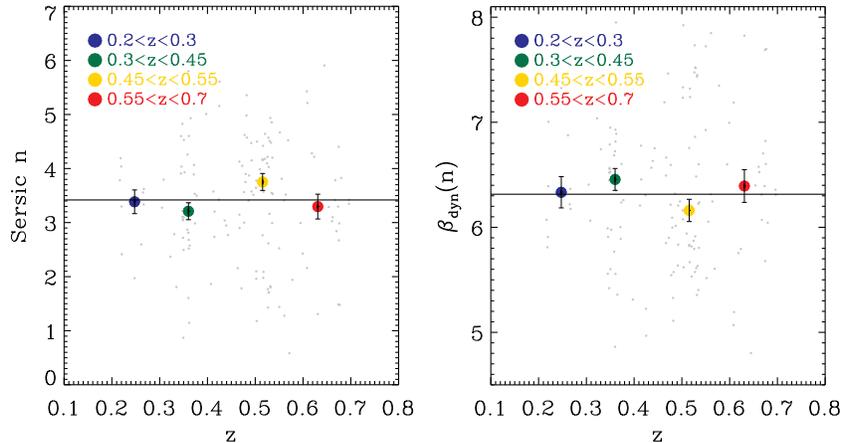


Figure 4. Sérsic indices (left panel) and β_{dyn} parameters (right panel; see Equation (3)) of the COSMOS/*HST* sample for various redshift bins as color-coded in Figure 2. Colored symbols with error bars are the medians. The continuous line is the median value. The total number of objects is 150. (A color version of this figure is available in the online journal.)

As we are mostly interested in the latter, the main conclusions of this work will not be affected. As an additional check we therefore compared the M_{dyn}/M_{\star} derived here with those measured for local galaxies from more sophisticated dynamical modeling in and find good agreement (see Section 4.3). Still, it should be emphasized that any change in dynamical mass found here reflects a change of dynamical mass within $1 R_e$.

3.4.1. Dependence on Structural Parameters

The appropriate value of β_{dyn} is actually a function of the Sérsic shape index n (Trujillo et al. 2004; Cappellari et al. 2006). Taylor et al. (2010) showed that dynamical masses and stellar masses correlate well when the structure of the galaxy is taken into account (see also Section 3.4.5). They find that dynamical masses estimated with the homology assumption exhibit residual trends with galaxy structure properties, so they introduce a structure-corrected dynamical mass adopting a constant β_{dyn} that is Sérsic index dependent (Bertin et al. 2002). Note that the virial mass estimator of Cappellari et al. (2006) (Equation (3)) has been calibrated on dynamical masses from Schwarzschild modeling where no assumption about homology is made.

For our sample of BOSS galaxies, we cannot make any statements in this respect since SDSS images do not have the necessary angular resolution to perform Sérsic fits. However, we can expect this effect to be negligible, as the BOSS galaxy sample is restricted to massive galaxies in a relatively narrow mass range (Maraston et al. 2013) and limited redshift range so that variations of the Sérsic index will be minimal. Moreover, the fact that our size calibration is based on Sérsic R_e from COSMOS, allows us to account for possible differences between de Vaucouleurs profiles and Sérsic profiles resulting in a “Sérsic-like” calibrated radii.

We verify this assumption with the COSMOS sub-sample for which Sérsic indices are available. The Zurich Structure & Morphology Catalog v1.0 also contains values of galaxy Sérsic index, n . This allows us, for this sub-sample, to account for the variation of the parameter β_{dyn} with n and encapsulate the effects of galaxy structure on M_{dyn} (by assuming a constant β_{dyn} for all galaxies, Equation (3) implicitly assumes that all galaxies are homologous).

We estimate $\beta_{\text{dyn}}(n)$ following the analytic expression between $\beta_{\text{dyn}}(n)$ and the Sérsic index (Equation (20) of Cappellari et al. 2006), which is theoretically derived for spherical and

isotropic models with a Sérsic profile for different values of n (Bertin et al. 2002; see also Taylor et al. 2010 for a discussion of its importance on the SDSS sample).

Figure 4, right panel, shows the dependence of the $\beta_{\text{dyn}}(n)$ parameter on redshift for each galaxy in the sub-sample (gray points). Colored circles are the median $\beta_{\text{dyn}}(n)$ for each redshift bin. We find that the median $\beta_{\text{dyn}}(n)$ is ~ 6.3 for all redshifts bins (see continuous line in Figure 4, right panel). This is larger than the local values of 5 generally adopted, and yields systematically higher masses by $\sim 20\%$. The reason is the relatively low Sérsic indices (between 3.38 and 3.30 at $z \sim 0.25$ or $z \sim 0.6$, as shown in Figure 4, left panel) for the COSMOS sample, compared to typical Sérsic indices for local galaxies.

The key point illustrated by Figure 4, however, is that both n and $\beta_{\text{dyn}}(n)$ do not evolve with redshift. As we focus in the redshift evolution and not absolute values for dynamical mass, the present study is not affected by a systematic offset in β_{dyn} . We will use M_{dyn} derived using a median $\beta_{\text{dyn}} = 6.3$, which is the median value β_{dyn} derived using the BOSS/COSMOS photometry.

3.4.2. Dependence on Aperture

The dynamical mass obtained using the virial mass estimator (see Section 3.4) is based on stellar kinematics within an aperture of 1 effective radius and scaled to *total* dynamical mass via Equation (3). This quantity is compared with the *total* stellar mass from Maraston et al. (2013) based on *cmodelmag* magnitudes. Hence both dynamical and stellar masses are *total* masses, which ensures a consistent comparison.

Still, the total dynamical mass is derived from observations within the effective radius, while the stellar mass comes from the total stellar light. We explore therefore the possible presence of a systematic effect from the different apertures in which kinematics and stellar populations have been measured. To this end we compare M_{\star} derived from *modelmag* (rescaled to *i*-band *cmodelmag*) and M_{\star} from aperture magnitudes within R_e (rescaled to *i*-band *cmodelmag*). This test is presented in Appendix A.

In brief, the difference between the two sets of masses is ~ 0.08 dex. The stellar masses measured from SED fitting within $1 R_e$ are higher by this amount, because of the higher M/L ratio within $1 R_e$. We emphasize, however, that this quantity is an overestimate of the true total mass. Nevertheless, it is

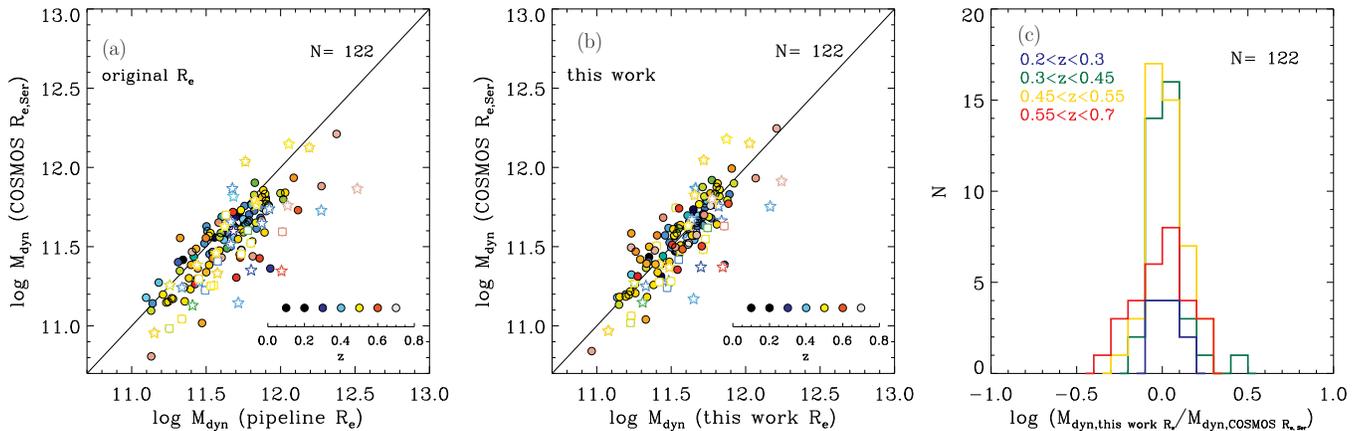


Figure 5. Left panel: dynamical masses derived using the original SDSS R_e assuming a constant $\beta_{\text{dyn}} = 6.3$ vs. dynamical masses derived using COSMOS $R_{e,\text{Ser}}$ adopting a variable β_{dyn} based on the Sérsic index. Symbols are as in Figure 1. Middle panel: same as left panel but using rescaled SDSS R_e adopted in the present study. Right panel: distribution of the ratio of dynamical masses from rescaled SDSS R_e and COSMOS $R_{e,\text{Ser}}$ for the four redshift bins of Figure 1. Legends in each panel give the number of galaxies used in the derivation of the size correction including objects discarded in the iterative 2σ clipping (multiple systems are not included in this number).

(A color version of this figure is available in the online journal.)

reassuring to verify that this systematic difference is relatively small. Most importantly, the offset is independent of redshift (see Appendix A). Hence, the science analysis of this work is not affected, because we study redshift dependence and do not focus on absolute ratios between dynamical and stellar masses. We also note that the dynamical to stellar mass ratio is always larger than 0.08 dex in our redshift range, hence M_\star does never exceed M_{dyn} ensuring physically meaningful solutions throughout.

3.4.3. Dependence on Rotation

The possible presence of unresolved rotation is another complication that could affect our mass estimates from Equation (3). van der Marel & van Dokkum (2007) have measured increased rotational support at $z \sim 0.5$ and argue that data at different redshifts can be affected by rotation, with a stronger impact on low- σ galaxies, which are more rotationally supported than galaxies at high σ . As our BOSS sample consists of massive galaxies in a relatively small mass range (Maraston et al. 2013), however, we expect this effect to be negligible.

3.4.4. Dependence on Galaxy Type

Finally, in deriving M_{dyn} with Equation (3), we implicitly assume that the measured value of σ_e is dominated by the bulge component. For late-type galaxies, we expect that the disc contribution to σ_e results in a broader distribution of M_{dyn} , since the σ_e may not represent the actual dynamical state of those galaxies which is dominated by rotation (see Section 3.4.3) as well as the β_{dyn} parameter we used could not be appropriate for late-type galaxies with low Sérsic index. As shown in Masters et al. (2011), however, the majority of BOSS galaxies (74% \pm 6%) have early-type morphology and the remaining later types are bulge dominated, hence this effect will be negligible. We tested this assumption by only considering early-type galaxies for the CMASS sample using the morphological cut ($g - i$) > 2.35 by Masters et al. (2011). We compared dynamical masses derived with COSMOS R_e and adopting β_{dyn} based on the Sérsic index and dynamical masses with R_e derived in this work and found a good agreement between early types and the full COSMOS/BOSS sub-samples, with a scatter around the one-to-one relation consistent within the errors (~ 0.14 dex).

3.4.5. Calibrated Virial Masses for the COSMOS Sub-sample

As an additional test, we compare our virial mass estimates based on the re-scaled effective radii with virial masses derived directly from the COSMOS effective radii, the result is shown in Figure 5. The left panel shows the comparison between virial masses derived using COSMOS R_e and the *uncorrected* SDSS R_e . As expected, there is a clear offset to higher virial masses from SDSS imaging because of the overestimation of galaxy radii.

The re-scaled radius of this work remedies this problem. The middle panel of Figure 5 shows the comparison between virial masses derived using COSMOS R_e , and adopting a variable β_{dyn} based on the Sérsic index (see Section 3.4.1) and the *corrected* SDSS R_e of the present work (by using a constant $\beta_{\text{dyn}} = 6.3$ as described in Section 3.4.1). Mass estimates agree well at all redshifts with a scatter of ~ 0.14 dex, which is well within the errors. The right panel presents the distribution of the logarithmic ratio between COSMOS and SDSS masses after correction. The distribution is symmetric around zero for all redshift bins with a maximum deviation of ~ 0.5 dex.

3.4.6. Random and Systematic Errors

The final errors in M_{dyn} are a combination of uncertainties in σ_e (which account for the aperture scaling of σ), R_e , the statistical uncertainties due to the rescaling factor of R_e , and β_{dyn} . This results in median random errors of ~ 0.15 dex depending on redshift (from 0.08 dex at $z = 0.25$ to 0.18 dex at $z = 0.65$). Based on Monte Carlo simulations, we estimate median systematic errors due to the size calibration (see Section 3.2.4) and uncertainty of $\beta_{\text{dyn}}(n)$ to be ~ 0.04 dex.

3.5. Local SDSS-II Early-type Galaxy Sample

We combine the SDSS-III/BOSS sample described above with a local sample of massive galaxies at $z \sim 0.1$ drawn from SDSS-II. The galaxy properties of this local sample are presented in the following sections.

3.5.1. Stellar Mass

Stellar masses and ages were estimated from the SED fitting of the u -, g -, r -, i -, z -band photometry following the same

prescription of Maraston et al. (2013) with passive templates (the LRG model by Maraston et al. 2009 mentioned earlier). We homogenize the stellar mass distribution by selecting a sub-sample that matches the mass distribution of the BOSS sample. We constructed this sub-sample using the stellar mass distribution in the lowest BOSS redshift bin ($0.2 < z < 0.3$) as reference. For each stellar mass bin, we randomly selected from the local galaxy distribution a number of galaxies equal to the number of galaxies in the low- z BOSS one. This cut on the local early-types population retains 12,089 galaxies. A discussion on the impact of the science analysis in this paper from this homogenization is given in Appendix E.

3.5.2. Size

We collect photometry and effective radii R_e from DR8 in which the correction for the sky over-subtraction of previous releases is already implemented (see discussion in Section 3.2.1), and no further correction to sizes (see Hyde & Bernardi 2009a, for details) has been applied. This is motivated by the fact that we selected galaxies at redshift $z < 0.2$ that are resolved in the SDSS imaging with $R_e > \text{FWHM}$ of the PSF (retaining 96% of the objects).

3.5.3. Stellar Velocity Dispersion

We collect redshift and stellar velocity dispersions from the DR7 catalogs. Thomas et al. (2013) show that their DR7 σ are consistent with SDSS pipeline σ at the few percent level (see their Figure 1). The median offset across all the stellar velocity dispersion is $\sim 1\%$. However, this offset increases toward high stellar velocity dispersions. We can quantify the correct offset to apply to DR7 σ looking at Figure 4 of Thomas et al. (2013) where their σ are compared to Bolton et al. (2012a) σ within BOSS, which is the relevant mass range.²⁴ The offset is 4%, which we correct for in the SDSS-II sample.

We further rescaled stellar velocity dispersions to the value at R_e , following the procedure described in Section 3.3, accounting for the fact that DR7 galaxies were observed with a $3''$ aperture. The variation in σ_e for the aperture correction in local SDSS galaxies is 2% (σ_e within R_e on average 2% smaller than the SDSS ones, and median ratio between aperture size and R_e is ~ 0.72).

3.5.4. Dynamical Mass

Dynamical mass is derived from stellar velocity dispersion and size in the same way as for the BOSS sample as described in Section 3.4. To ensure internal consistency, we use the same redshift independent parameter $\beta_{\text{dyn}} = 6.3$ as for the BOSS sample derived from the BOSS/COSMOS photometry (see Section 3.4.1).

3.6. Correction for Progenitor Bias

BOSS target selection was designed to obtain a nearly uniform stellar mass distribution across the redshift range $0.2 \leq z \leq 0.7$. Still, the sample needs to be corrected for effects from progenitor biases (e.g., Valentinuzzi et al. 2010b; Saglia et al. 2010; Cimatti et al. 2012 and references therein), as higher- z galaxies in the sample are not necessarily progenitors of the lower- z galaxies in the sample (see also Tojeiro et al. 2012).

²⁴ Bolton et al. (2012a) is the same code that produced the SDSS σ .

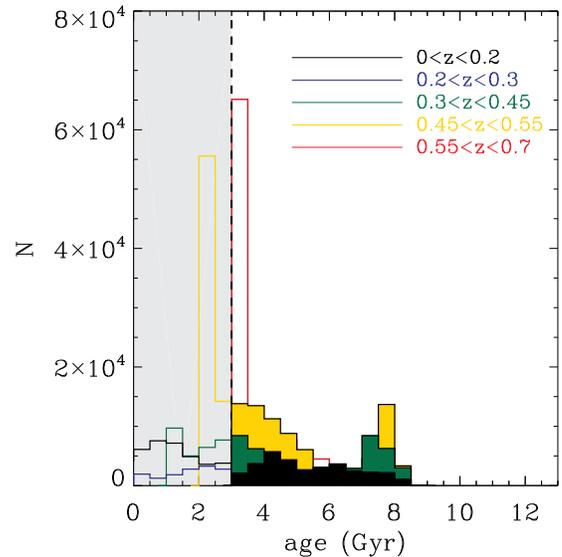


Figure 6. Histograms of the ages for different redshift bins, evolved to the highest redshift bin by subtracting the look-back time, which are used to apply the progenitor-bias cut (see the text for details). Galaxies in the shaded region, i.e., with an age below 3 Gyr, have been discarded.

(A color version of this figure is available in the online journal.)

To correct for the progenitor bias, we compare—in each redshift bin—the galaxy ages from Maraston et al. (2013; one of the products of the SED fit; see Section 3.1) and remove those galaxies from the low- z sample whose ages (evolved to the highest redshift bin by subtracting the look-back time) would be lower than a given age threshold which is the time needed for a typical galaxy to become passive. For each redshift bin, we select galaxies such that their age follows

$$t_g(z) - (t_u(z) - t_u(z = 0.65)) > 3 \text{ Gyr}, \quad (4)$$

where t_g is the age of a galaxy at a give redshift, t_u is the age of the universe at the same redshift, and $t_u(z = 0.65)$ is the age of the universe at the median redshift of the highest-redshift bin. Histograms of the evolved ages for different redshift bins are shown in Figure 6. Galaxies in the shaded region have been discarded. As the age threshold we chose 3 Gyr, adopting the age limit used in Maraston et al. (2013)²⁵ for calculating stellar masses (see their Section 3.1 for discussion). This threshold is only slightly larger than the 1.5 Gyr suggested by van Dokkum & Franx (2001).

We considered the highest-redshift bin as a reference and we evolved all other redshift bins including the local SDSS early-type sample. By discarding galaxies with age < 3 Gyr, we retain 268,938 galaxies, which correspond to the $\sim 65\%$ of the initial local and BOSS samples as shown in Figure 6.

We obtain similar results using the tighter selection criteria described in Cimatti et al. (2012), which select in each redshift bin the galaxies with ages within $\pm 1\sigma$ of the age distribution for each redshift bin accounting for the cosmic time elapsed from one bin to the other. This selection also discards objects at older ages and provides a sample size that is $\sim 54\%$ of the initial one.

Poggianti et al. (2013) found that galaxy sizes are correlated to luminosity-weighted ages such that older galaxies will show

²⁵ Maraston et al. (2013) set a minimum age of 3 Gyr for the mass calculation using the passive template in order to minimize the chance to underestimate the mass by underestimate the galaxy age. This age limit translates into the assumption of a high-formation epoch for the massive and passive galaxies in CMASS.

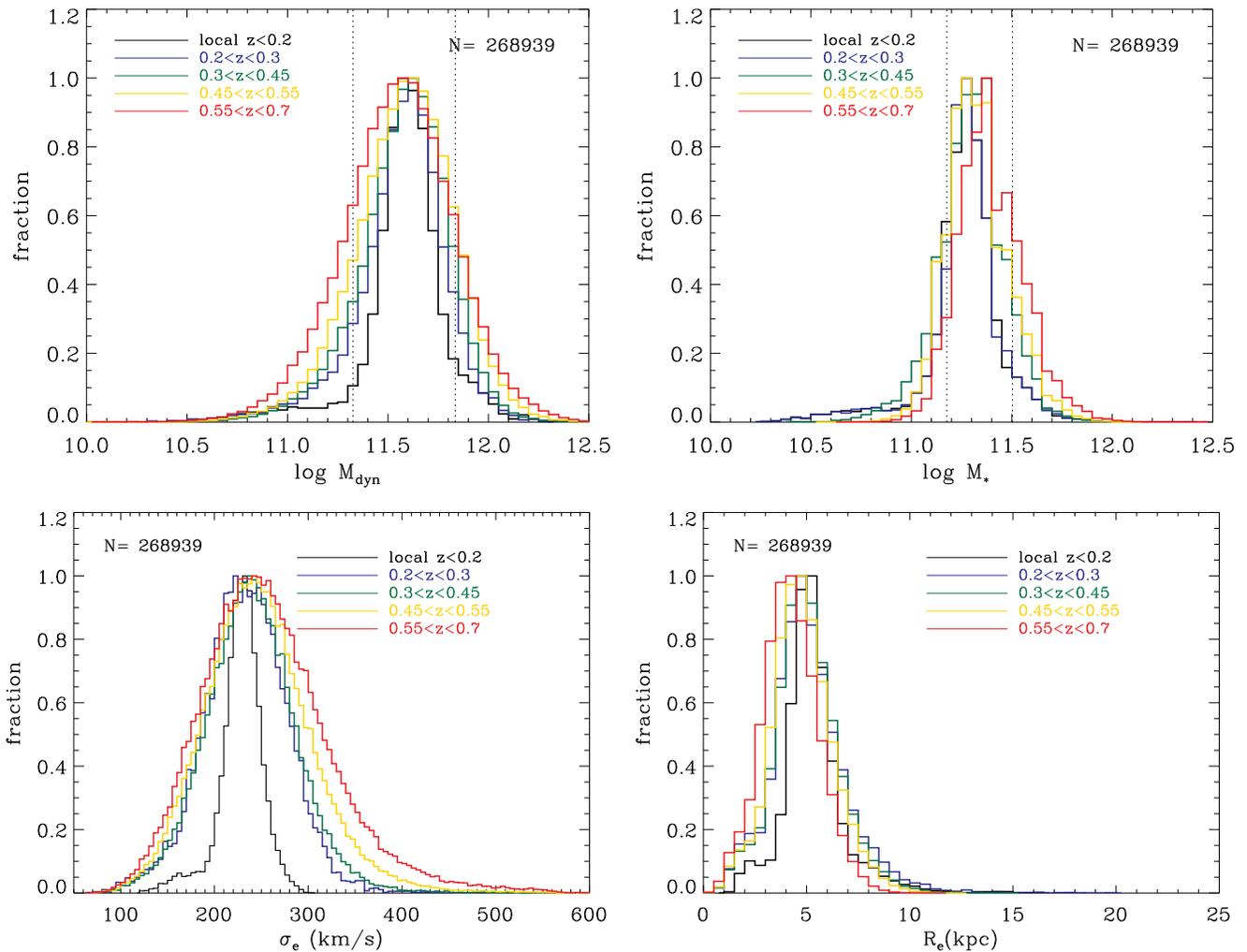


Figure 7. Top panels: distribution of M_{dyn} (with $\beta_{\text{dyn}} = 6.3$, left panel) and M_* (right panel) for various redshift bins, normalized to the peak value in each bin. The BOSS mass distributions are fairly uniform over the redshift range under analysis (see also Figure 10 in Maraston et al. 2013). Local early-types from SDSS-II are selected to have the same stellar mass distribution of the lowest BOSS redshift bin. Dotted black lines indicate the $\pm 1\sigma$ of the mass distributions adopted for the present analysis. Bottom left and right panels: distributions in stellar velocity dispersion distribution and effective radius. The progenitor-bias correction has been applied in all cases.

(A color version of this figure is available in the online journal.)

a stronger size evolution, with a stronger effect in clusters than in the field. Our progenitor bias correction minimizes that effect.

The distributions of M_* , M_{dyn} , σ_e , and R_e of the final sample after correction for progenitor bias are shown in Figure 7 for various redshift bins. The typical median stellar mass is around $\log M_* \sim 11.28$ dex, the median $R_e \sim 5.2$ kpc, and $\sigma_e \sim 231 \text{ km s}^{-1}$.

To study the effect of the progenitor bias correction on the redshift evolution of these quantities, we have performed a re-analysis for a sample without progenitor bias correction presented in Appendix D. It can be seen that generally results are consistent. Most importantly, the evolution of M_{dyn}/M_* is fairly stable against the progenitor-bias correction, hence the main conclusions of these paper do not critically depend on the progenitor bias correction.

4. RESULTS

In this section, we present the redshift evolution of the galaxy parameters effective radius, velocity dispersion, and dynamical to stellar mass ratio for our final sample of 256,849 SDSS-III/BOSS galaxies and 12,089 local early-type SDSS-II

galaxies with a typical stellar mass of $\sim 2.2 \times 10^{11} M_\odot$ and a typical dynamical mass of $\sim 3.9 \times 10^{11} M_\odot$. The results are presented in Figure 8, where we plot the galaxy parameters effective radius (left panels), stellar velocity dispersion (middle panels), and dynamical to stellar mass ratio (right panels) as functions of redshift. Shaded regions and contours indicate the number density of galaxies (10 equally-spaced density levels showing the percentage of galaxies compared to the peak value of each plot), and colored circles are the mean for each redshift bin. Fixed intervals in stellar mass and dynamical mass are considered in the top and bottom panels, respectively. They were selected to be within $\pm 1\sigma$ of the mass distributions of Figure 7. This allows us to keep a large number of galaxies with similar mass (186,269 and 189,613 galaxies for M_* and M_{dyn} selection for the full local and BOSS sample, respectively) without being affected by selection effects as a function of z . A finer division in both M_* and M_{dyn} would not change our results.

The solid line in Figure 8 is a linear fit to the relation, whereas the dashed line is a fit to the zero point at constant zero slope and the black dotted line is the linear fit to a sample without size correction. The fit parameters are summarized in Table 2.

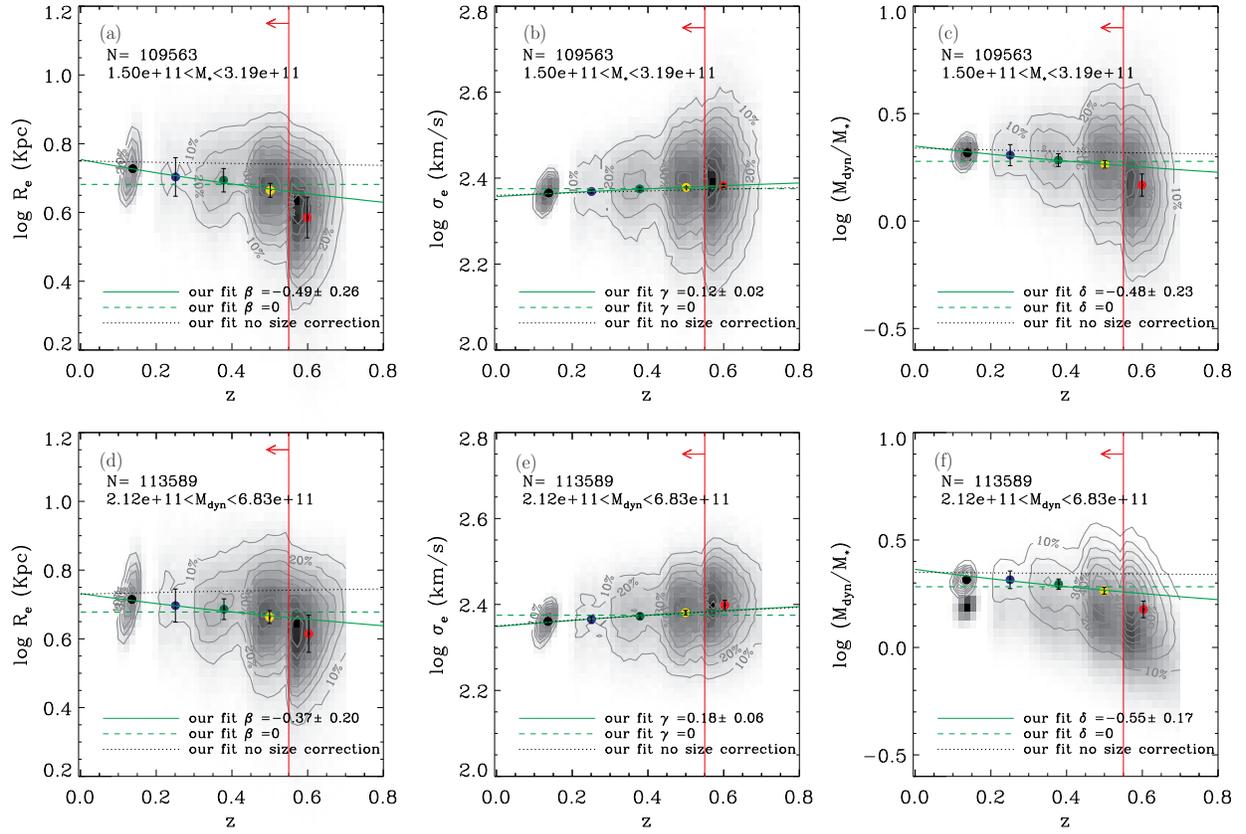


Figure 8. Left panels: effective radius R_e as a function of redshift. Middle panels: stellar velocity dispersion σ_e as a function of redshift. Right panels: ratio between dynamical and stellar mass M_{dyn}/M_* as a function of redshift. Top and bottom panels are for galaxies selected using M_* and M_{dyn} , respectively (within $\pm 1\sigma$ of the mass distributions, total number of galaxies given by the labels). The shaded contour region indicates the full sample after correction for progenitor bias. Contours show 10 equally-spaced density levels showing the percentage of galaxies compared to the peak value of each plot. The colored filled circles are the median values in four redshift bins. The green solid line is a linear fit, the green dashed line is a fit with zero slope for comparison. The black dotted line is a linear fit to the sample if no size correction is applied. Error bars of the median points indicate the 1σ uncertainty derived from Monte Carlo simulations. The red continuous lines and arrows indicate the range where we fit our data.

(A color version of this figure is available in the online journal.)

Table 2
Fitting Parameters for the Redshift Evolution of Galaxy
Parameters between $0.1 \leq z \leq 0.55$

Parameter	M_*		M_{dyn}	
	Slope	Zero Point	Slope	Zero Point
R_e	-0.49 ± 0.26	0.76 ± 0.04	-0.37 ± 0.20	0.73 ± 0.03
σ_e	0.12 ± 0.02	2.36 ± 0.01	0.18 ± 0.06	2.35 ± 0.01
M_{dyn}/M_*	-0.48 ± 0.23	0.35 ± 0.03	-0.55 ± 0.17	0.36 ± 0.02

Notes. Uncertainties on each parameter are 1σ errors derived from Monte Carlo simulations. The relation we fitted for R_e is $\log R_e = \log R_{e,0} + \beta(1+z)$, for σ_e is $\log \sigma_e = \log \sigma_{e,0} + \gamma(1+z)$, and for M_{dyn}/M_* is $\log(M_{\text{dyn}}/M_*) = \log(M_{\text{dyn}}/M_*)_0 + \delta(1+z)$.

We fit relationships of the form $X \propto (1+z)^{\text{slope}}$ to all data, but the result does not change significantly by fitting the means. We do not account for galaxies in the last redshift bin at $z > 0.55$ in the fit because of the larger uncertainty of the radius calibration (see Section 3.2.3). The best-fitting values of zero-point, slope, and their associated errors are derived by performing least-squares linear regressions using the MPFIT package. We additionally consider the case where we only fit the zero-point (assuming a zero slope) to test the significance of the derived slopes (dashed lines). We also assess the latter with

a comparison of the χ^2 values of the two fits for free and fixed slope, accounting for the number of degrees of freedom, using the F -test statistics.

Figure 9 is a reproduction of Figure 8 in which the predictions of simulations are shown for comparison. Solid lines in Figure 9 are model predictions of Oser et al. (2012), Nipoti et al. (2012) (here we list a couple of models with different stellar-to-halo-mass prescriptions as a function of redshift that those author presented in their work), Hopkins et al. (2009), and Khochfar & Silk (2006) for the evolution of galaxy size and velocity dispersion. The predictions of the redshift evolution of M_{dyn}/M_* are from Hopkins et al. (2009).

As discussed in the previous sections, the major sources for random and systematic errors are the size correction (Section 3.2) and the calculation of dynamical mass through the virial estimator (Section 3.4). To assess random and systematic errors in the redshift evolution of the galaxy parameters, we perform Monte Carlo simulations perturbing the slope a , the intercept b of the size correction, as well as the structural dependent quantity $\beta_{\text{dyn}}(n)$ within their errors. For each redshift bin, we produced distributions of a , b , and β_{dyn} generating random numbers from their errors and assuming normal distributions. For each of the 200 realizations we then derived the mean slope and the 68% confidence intervals for the evolution of R_e , σ_e , M_{dyn}/M_* as a function of redshift. The error bars in Figures 8

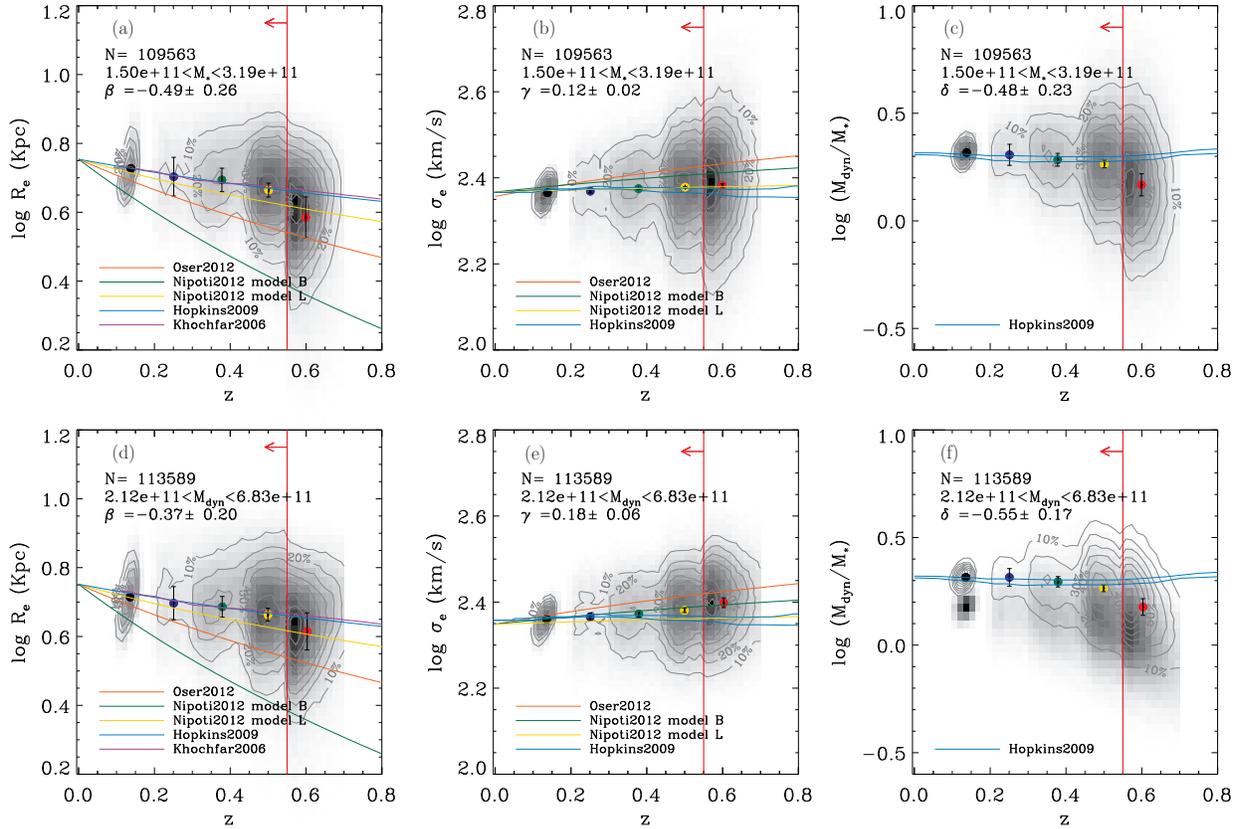


Figure 9. Same as Figure 8, but with predictions from theoretical models overplotted as solid colored lines for comparison (see labels for references).

(A color version of this figure is available in the online journal.)

and 9 are estimated through these simulations hence include both random and systematic errors.

4.1. Evolution of Galaxy Size

The left panel of Figure 8 shows that galaxy radius decreases with increasing redshift for both choices of mass estimator (top and bottom panels) at 1.5σ significance. The F -test between the fits with fixed and free slope yields a probability $<25\%$ of the null hypothesis being true (no redshift evolution) for both M_* and M_{dyn} selected samples, which supports the significance of the slope derived here. Qualitatively similar results are obtained when only using the BOSS sample, although uncertainties are larger and the significance reduced (Appendix C). The size evolution found in the present work are consistent within the errors with previous determinations in the literature, which are mostly based on data at higher redshifts (e.g., Trujillo et al. 2006a; van Dokkum et al. 2008; Cimatti et al. 2008; Saracco et al. 2009), but in particular with Saglia et al. (2010), who studied a similar z range. This agreement further validates the size correction applied here. Without the latter, we would not detect significant evolution of galaxy sizes (dotted lines in Figure 8) in clear contradiction to findings in the literature.

We note again that we did not account for the slightly different mapped rest-frame wavelengths using radii from observed i -band images across all redshifts. This approach is conservative as even smaller sizes would be obtained from the rest-frame bluer images at higher redshift (Bernardi et al. 2003; Hyde & Bernardi 2009a) with the net effect that we tend to slightly underestimate the size evolution.

In Figure 9 (left panels), we show the comparison of our results with simulations (solid lines), which show a very wide range of predictions for the slope β . The evolution we find is consistent with or slightly milder than the predictions from semi-analytical models or hydrodynamical simulations (Khochfar & Silk 2006; Naab et al. 2009; Hopkins et al. 2009; Nipoti et al. 2012; Oser et al. 2012). However, recent work on size evolution suggests that the size evolution at $z < 1$ is much shallower than at high redshifts (Newman et al. 2012; Cimatti et al. 2012; Nipoti et al. 2012), which could explain why we find a milder evolution of the effective radius. This is reinforced by comparison with the predictions of López-Sanjuán et al. (2012), who studied close pairs using massive galaxies in COSMOS up to $z \sim 1$ and measured the merger fraction and rate from both minor and major mergers. Their models predict a size evolution due to both major and minor mergers, which is in excellent agreement with our results.

4.2. Evolution of Stellar Velocity Dispersion

The middle panels of Figure 8 show the evolution of σ_e with redshift. We detect a mild but significant evolution of stellar velocity dispersion, in the sense that σ_e increases with increasing redshift at $>2\sigma$ significance. Again, the F -test between the fits with fixed and free slope yields a probability $<1\%$ or $<2\%$ of the null hypothesis being true (no redshift evolution) for M_* and M_{dyn} selected samples, respectively, which supports the significance of the slope derived here. As for the size evolution, qualitatively similar results are obtained when only using the BOSS sample again with somewhat larger uncertainties and slightly reduced significance (Appendix C).

Our findings are consistent with previous results from the literature (Cenarro & Trujillo 2009; van Dokkum et al. 2009; Saglia et al. 2010; van de Sande et al. 2011, 2013), although the evolution detected here is somewhat milder, possibly because of the relatively small redshift range mapped in the present work. Indeed, studying a similar redshift range, Saglia et al. (2010) find that, depending on the selection criteria and accounting for progenitor bias, the slope γ ranges from 0.59 ± 0.10 to 0.19 ± 0.10 , which is consistent with our results.

In Figure 9 (middle panels), we compare our results with the large range of predictions of γ coming from simulations. The evolution is consistent with the predictions from the models by Oser et al. (2012) if selected by dynamical mass (middle bottom panel in Figure 9). A milder evolution of σ_e , however is predicted by the hydrodynamical simulations of Hopkins et al. (2009) and N -body simulations of Nipoti et al. (2012) which are consistent with the middle top panel in Figure 9. An excellent agreement is found comparing our results with the predictions on the σ evolution from the work by López-Sanjuan et al. (2012) discussed in the previous section.

Hopkins et al. (2009) suggest that velocity dispersions do not evolve significantly with redshift for the redshift range probed here; they find a mild evolution at $z > 1$, which they explain with velocity dispersions being set by the dark matter halos that evolve more weakly compared to R_e . This absence of evolution in our redshift range is in tension with the observational result presented here.

4.3. Evolution of the Dynamical to Stellar Mass Ratio

The right panels of Figure 8 display the evolution of the dynamical to stellar mass ratio M_{dyn}/M_* with redshift. This ratio decreases with increasing redshift at $>2\sigma$ significance. Again, the F -test between the fits with fixed and free slope yields a probability $<2.5\%$ and $<1\%$ of the null hypothesis being true, for M_* and M_{dyn} selected samples, respectively, which supports the significance of the slope derived here. The decrease is driven by the decrease in size, and not balanced by the very mild increase in stellar velocity dispersion. The slopes for the evolution are consistent within the errors whether we select our sample by stellar or dynamical mass. As for the evolution of galaxy size and velocity dispersion qualitatively similar results are obtained when only using the BOSS sample, again with somewhat larger uncertainties and slightly reduced significance (Appendix C).

Also in Appendix C, we discuss the effect of a redshift-variable parameter $\beta_{\text{dyn}}(n)$ for the BOSS sample. We ran additional simulations with redshift-dependent structural parameter n and hence a redshift-dependent virial constant $\beta_{\text{dyn}}(n)$ based on the COSMOS/*HST* measurements. In brief, we find that the results of this paper are not affected. This ought to be expected as the redshift evolution of n (and hence $\beta_{\text{dyn}}(n)$) is mild as discussed in Section 3.4.1.

Our finding of a decreasing M_{dyn}/M_* ratio with increasing redshift is consistent with the evolution of R_e and σ_e from Saglia et al. (2010), resulting in a similar trend of decreasing M_{dyn}/M_* with redshift at a 2σ level. We searched for systematic effects by checking the evolution of M_{dyn} and M_* separately. For fixed M_* , M_{dyn} decreases with z ($\delta = -0.23 \pm 0.12$) whereas for fixed M_{dyn} , M_* increases with z ($\delta = 0.85 \pm 0.11$).

A change in dynamical to stellar mass ratio can have several physical explanations. In general, the effects of varying dark matter fraction and change in the inferred stellar mass due to a variable IMF are highly degenerate, and it is notoriously difficult

to distinguish the two. In the present study, we use an approach in which redshift evolution is added as a further constraint. As we are probing a well-selected, passively evolving galaxy sample consisting of low- z massive galaxies and their high- z progenitors, any variation in stellar population property would be minor (in case of galaxy mergers, for instance, the variation of the effective IMF would be small). As a consequence, the decrease of M_{dyn}/M_* with redshift is most plausibly caused by a decrease of dark matter fraction. We emphasize again that we are probing a variation of stellar kinematics within the effective radius, hence a possible change of dark matter fraction within $1 R_e$, even though total masses are compared. In other words, our results imply that the dark matter fraction in massive galaxies within the half-light radius increases with cosmic time.

4.3.1. Comparison with Local and $z \sim 0.8 M_{\text{dyn}}$

It is worth noting that the mean ratio between dynamical and stellar mass is larger than one at all redshifts. This point is crucial, as a smaller dynamical than stellar mass would be unphysical. As a key consistency check, we compare our M_{dyn}/M_* values with derivations for local galaxies based on sophisticated dynamical modeling by Thomas et al. (2011) as opposed to the simple virial mass estimator adopted here. Thomas et al. (2011) derive dynamical to stellar mass ratios of 1.8 (assuming a Kroupa 2001 IMF) for a sample of early-type galaxies in the Coma cluster. The value derived in the present work for the lowest redshift bin, $z \sim 0.1$, is ~ 2.1 (also based on Kroupa 2001 IMF), is well consistent with this value, as well as with other published values for the SDSS sample (~ 1.7 ; Taylor et al. 2010).

Recent work of Shetty & Cappellari (2014) found that galaxies at $z \sim 0.7-0.9$, with stellar mass $M_* > 10^{11} M_\odot$ and stellar velocity dispersion $\sim 200 \text{ km s}^{-1}$, have an average normalization of the IMF consistent with a Salpeter slope, similarly to recent findings in the local universe (e.g., Cappellari et al. 2012). In our work, we cannot constrain the actual normalization of the IMF, but we note that 7% of our BOSS galaxies with stellar velocity dispersion between $200 < \sigma_e < 280 \text{ km s}^{-1}$ and errors on the stellar velocity dispersion smaller than the typical cut we use in our analysis ($<10\%$), would have unphysical M_{dyn}/M_* ratio by using a Salpeter IMF. Those galaxies also have a smaller average size, $\sim 3.5 \text{ kpc}$, compared to the typical $\sim 4.6 \text{ kpc}$ of galaxies with a physical M_{dyn}/M_* ratio. By using a Kroupa IMF, only the 0.5% of BOSS galaxies with $200 < \sigma_e < 280 \text{ km s}^{-1}$ and errors on $\sigma_e < 10\%$, have an unphysical M_{dyn}/M_* ratio (those galaxies have also an average size of $\sim 2 \text{ kpc}$).

4.3.2. Comparison with Simulations

The right panels of Figure 9 show the comparison between our results and simulations by Hopkins et al. (2009) for galaxies with $M_* \sim 10^{11} M_\odot$ (blue lines). The latter predict almost no evolution of M_{dyn}/M_* with redshift for galaxies at $M_* \sim 10^{11} M_\odot$ in our redshift range, in tension with our observational result. Hopkins et al. (2009) predict that M_{dyn}/M_* decreases with increasing redshift beyond $z \sim 1$. The recent galaxy formation simulations by Hilz et al. (2012, 2013), instead, are in better agreement with our observations. The authors find that galaxy sizes grow significantly faster and the profile shapes change more rapidly for minor mergers of galaxies embedded in dark matter halos than for major mergers. Moreover, the increase in stellar mass is much smaller for minor mergers than for major mergers. This growth is accompanied by an increase of the dark

matter fraction within the half-mass radius, driven by the strong size increase probing larger, dark matter dominated regions (Hilz et al. 2013). In this scenario, the dark matter fraction in the center of a galaxy is expected to increase with cosmic time, in agreement with the observational result found in the present study. As shown in Hilz et al. (2013), major mergers could also result in an evolution of M_{dyn}/M_{\star} , although by a smaller amount (25%), and they would change substantially the internal structure of the galaxy. We caution that our data do not constrain any difference between minor and major merger, we only study the relative evolution of galaxy properties and not absolute quantities. Minor mergers could explain our results but this does not exclude that major mergers play a role as well.

5. DISCUSSION

In the past years, a number of investigations have been dedicated to studying the dependence of dynamical to stellar mass ratios M_{dyn}/M_{\star} with galaxy mass in the local universe. There is a clear concordance that M_{dyn}/M_{\star} increases with galaxy mass. The origin of this trend remains controversial, however. It is still under debate whether this phenomenon is driven by dark matter fractions, variations of the IMF, non-homology of early-type systems or adiabatic contraction (Cappellari et al. 2006, 2012; Hyde & Bernardi 2009a; Treu et al. 2010; Auger et al. 2010b; Grillo 2010; Napolitano et al. 2010; Schulz et al. 2010; Dutton et al. 2011, 2012, 2013; Thomas et al. 2011; Conroy & van Dokkum 2012; van Dokkum & Conroy 2012; Wegner et al. 2012; Conroy et al. 2013).

In this paper, we study the evolution of the dynamical to stellar mass ratio of massive galaxies as a function of cosmic time. The extra dimension added with look-back time helps to break some of the degeneracies plaguing local studies, because we analyze a passively evolving galaxy population in a very small mass range (see Figure 7). In this case variations of the effective IMF due to mergers would be minor. We find that the dynamical to stellar mass ratio in massive galaxies of $M_{\star} \sim 2 \times 10^{11} M_{\odot}$ decreases with increasing redshift at $>2\sigma$ significance over the redshift range $0.1 < z < 0.55$.

5.1. Comparison with High-redshift Literature Data

The relatively modest evolution of M_{dyn}/M_{\star} over the past seven billion years found here is well in line with other studies in the literature generally probing higher redshift and larger look-back times. The SDSS-III/BOSS data serve well in bridging galaxy properties from the distant with the local universe. In this section we will put those two data sets together comparing our results directly with the data at high redshift.

We collect public catalogs of structural parameters, stellar masses and stellar velocity dispersions from the EDisCS survey described in Saglia et al. (2010), for a sample of 154 cluster and field galaxies (41 field galaxies and 113 cluster galaxies) at median redshift $z \sim 0.7$. We derive dynamical masses as described in Section 3.4 using a variable $\beta_{\text{dyn}}(n)$ derived from EDisCS Sérsic indices and Equation (20) of Cappellari et al. (2006). We rescale sizes in kiloparsecs of Tables 1 and 2 of Saglia et al. (2010) to our cosmology and stellar velocity dispersions are rescaled to R_e using Equation (2). Saglia et al. 2010 stellar masses, derived using Bruzual & Charlot (2003) models and a diet-Salpeter IMF (Bell & de Jong 2001), are rescaled to a common Kroupa IMF (to match the IMF used for our local SDSS and BOSS sample), using a -0.05 dex offset

based on Table 2 of Bernardi et al. (2010). We select galaxies with $\log M/M_{\odot} > 11$ resulting in 77 objects.

We collect $z > 1$ data from van de Sande et al. (2013). These authors presented five new kinematic measurements of galaxies at $z > 1.5$ and compiled a catalog of previous data in the range of $0.8 < z < 2.18$ (van der Wel et al. 2008; van Dokkum et al. 2009; Newman et al. 2010; Onodera et al. 2012; Toft et al. 2012; Bezanson et al. 2013b), for a total of 73 galaxies, 46 of which have $\log M/M_{\odot} > 11$. The dynamical masses in van de Sande et al. (2013) were derived by using procedures similar to those described in Section 3.4, accounting for a variable $\beta_{\text{dyn}}(n)$, hence we only rescale them to our cosmology. Stellar masses, derived using Bruzual & Charlot (2003) models and a Chabrier (2003) IMF, were rescaled to a common Kroupa IMF using a $+0.05$ dex offset based on Table 2 of Bernardi et al. (2010).

5.1.1. Stellar Masses

Stellar masses derived with different population models may be different because of the different assumptions of stellar evolution in the models. Moreover, other assumptions regarding the star formation history, dust reddening and the assumed IMF all affect the final value of M_{\star} .

The variation in M_{\star} is quantified in Pforr et al. (2012) as a function of population model, using Maraston (2005) and Bruzual & Charlot (2003) models, and as a function of the star formation history and IMF assumed in the models. We use their results for obtaining a homogeneous sample of M_{\star} .

As most of the galaxies in the $z > 0.7$ sample studied here appear to be passive (van de Sande et al. 2013), we choose offsets from Table B4 of Pforr et al. (2012) derived for mock passive galaxies at $z = 2$. As a fitting setup, we select the “wide BC03” with reddening included as adopted in the $z > 0.7$ literature stellar masses, which gives an offset to Maraston (2005) based stellar masses of 0.13 dex. We note that the star-forming mocks show the same offset (0.12 dex) for the “wide BC03” fitting setup, which is important as some of the galaxies in the sample might not be passive (see, for instance, the Bezanson et al. 2013b sample). We decrease the stellar masses of the $z > 0.7$ sample by this amount. This offset is consistent with differences in stellar mass due to stellar population models found for BOSS galaxies (see Appendix of Maraston et al. 2013) and for galaxies in COSMOS (Ilbert et al. 2010).

It should be noted, however, that stellar masses also depend on the star formation history adopted for the SED fitting. Stellar masses of the $z > 0.7$ sample were obtained assuming an exponentially decaying star formation history. However, galaxies at those redshift may be better modeled with an exponentially increasing star formation history (Maraston et al. 2010), which would give higher stellar masses by ~ 0.2 dex compensating the offset applied here. Ideally the full sample should be modeled self-consistently, but the photometry of the $z > 0.7$ sample is not available to us. We will therefore discuss final results based on both with and without the offset of 0.13 dex in stellar mass.

5.1.2. Evolution of M_{dyn}/M_{\star}

Figure 10 presents the evolution of the dynamical to stellar mass ratio as a function of redshift for the redshift interval $0.08 < z < 2.18$. The SDSS-II and SDSS-III/BOSS data of the present study is combined with the high- z samples discussed in the previous section. The 0.13 dex offset between the stellar masses of the low- and high- z samples has been applied.

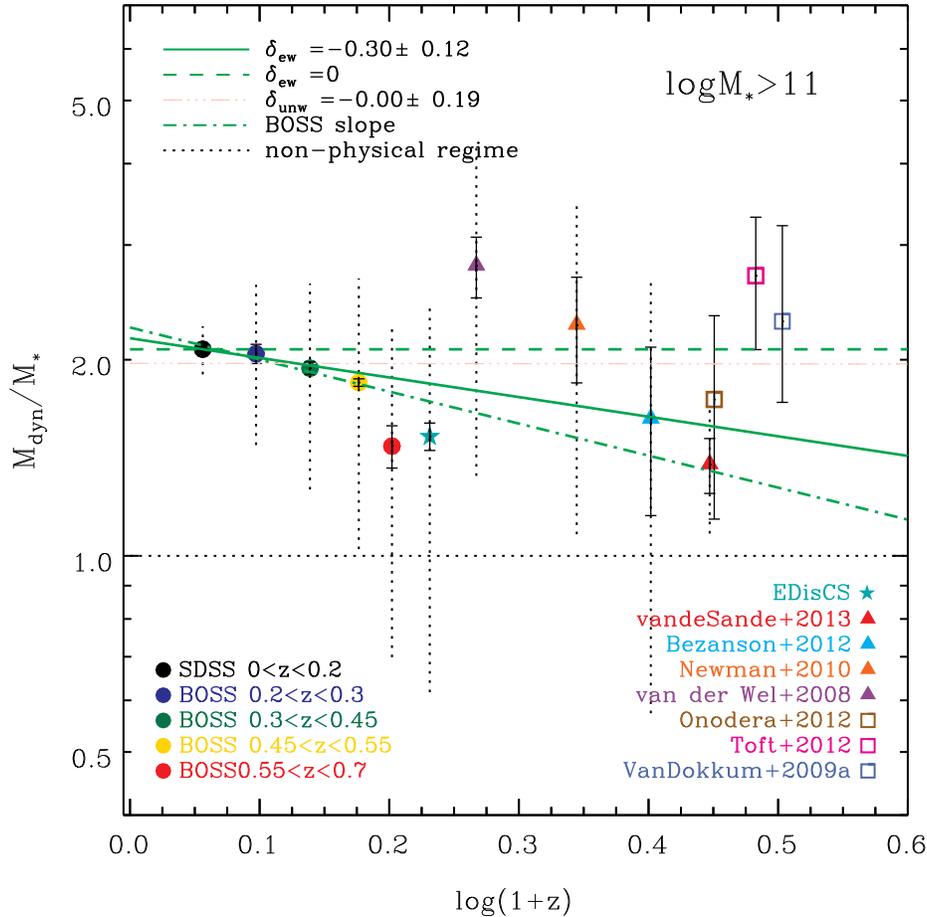


Figure 10. Ratio between dynamical and stellar mass as a function of redshift extending up to $z \sim 2$ for galaxies with $\log M_*/M_\odot > 11$. Data samples from the literature with $z > 0.7$ (see labels in the plot) have been added to the sample presented in Figure 8. The filled colored circles are the median M_{dyn}/M_* ratios for the redshift bins of Figure 8. The filled star and triangles are the median M_{dyn}/M_* ratios for each sample from the literature. Open squares are individual measurements for publications of only one object. The green solid line is the error-weighted linear fit to the median values plotted covering the full range in redshift $0.08 < z < 2.18$, whereas the pink line is the unweighted linear fit to the same points. The green dashed line is a fit with zero slope for comparison. The green dot-dashed line is the linear fit to our sample from Figure 8. The black dotted horizontal line indicates the non-physical regime of $M_{\text{dyn}}/M_* < 1$ (below the line). Black dotted vertical lines indicate the standard deviation of the distribution at the given redshift.

(A color version of this figure is available in the online journal.)

The symbols plotted are the median values of both M_{dyn}/M_* and z . Error bars are standard errors, while the dotted lines indicate the standard deviation of the distribution at the given redshift.

The high- z sample is consistent with the trend of decreasing M_{dyn}/M_* , even though the scatter at $z > 1$ is large. The BOSS data at intermediate redshifts clearly drives this relationship, because of the large scatter in the data at high redshift. By fitting the data over the full redshift range of $0.08 < z < 2.18$ as shown in Figure 10 and including the error bars in the fit, we find $M_{\text{dyn}}/M_* \propto (1+z)^{-0.30 \pm 0.12}$ (where δ_{ew} is the slope, and “ew” stands for error-weighted fit). This slope is slightly shallower but well consistent with the value derived in this work from the SDSS-II and SDSS-III/BOSS data alone (dot-dashed line; see Figure 8). Most importantly, the statistical significance for the presence of a negative slope is $>2\sigma$ also in this case. This further reinforces the evidence for a decrease in M_{dyn}/M_* with increasing redshift.

Error weighting the fit could potentially bias our results toward the BOSS sample, where the statistic is much larger and errors on the mean values are smaller. Therefore, we repeated the same procedure above using an equal weighting for all the points. This is shown by the pink triple-dot-dashed line

in Figure 10 (where δ_{unw} is the slope and “unw” stands for unweighted fit). In this case we find almost no evolution as we could easily expect. Given the large uncertainties and the low statistic in the high- z sample, applying a different weight to the BOSS sample would be preferable. A larger sample of high-redshift data would be helpful to constrain the evolution M_{dyn}/M_* ratio over a redshift range wider than that of BOSS.

Figure 11 shows the case in which the offset to the stellar masses has not been applied. As the correction implied a decrease of stellar masses in the high- z sample, the decrease of M_{dyn}/M_* with increasing redshift becomes even steeper and the statistical significance increases to $>4\sigma$.

5.2. Decreasing Dark Matter Fraction due to Size Growth

The increase of M_{dyn} with cosmic time is most plausibly caused by an increase of dark matter fraction within the effective radius. This increase can be well understood through size growth, which causes an increase of the dark matter fraction within an increasing effective radius (van de Sande et al. 2013). Indeed simulations show that the addition of stars in the outskirts of galaxies following the minor merger scenario can lead to an increased measured dark matter fraction by $\sim 80\%$ (Johansson

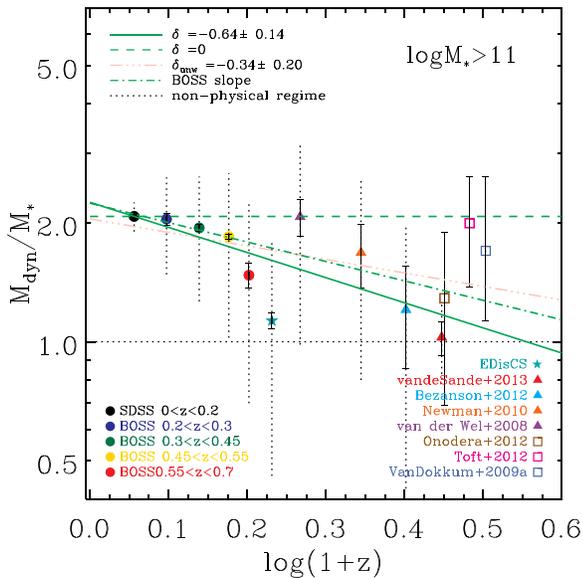


Figure 11. Same as Figure 10, but without correction for offset in stellar masses due to differences in stellar population modeling (see the text for details).

(A color version of this figure is available in the online journal.)

et al. 2012; Hilz et al. 2012) because more area with larger dark matter fraction is also included within R_e (Hilz et al. 2013; Hopkins et al. 2009). Also, Toft et al. (2012) studying galaxies at $z \sim 2$ with available kinematics suggest that the low dark matter fraction of galaxies at $z \sim 2$ is in favor of the merger scenarios which can redistribute dark matter within R_e (Boylan-Kolchin et al. 2005; Oser et al. 2012).

The steady increase of the dark matter fraction in the centers of massive galaxies with time further implies that massive galaxies in the local universe must contain some dark matter within their half-light radii, even if they are baryon dominated. This is consistent with recent dynamical modeling of nearby galaxies implying dark matter fractions of 16%–28% (Thomas et al. 2011; Cappellari et al. 2013b) as well as simulations predicting dark matter fractions of 18%–38% (Naab et al. 2007, and references therein).

6. CONCLUSIONS

We study the redshift evolution of the dynamical properties of $\sim 180,000$ galaxies from the SDSS-III/BOSS survey. We examine the redshift evolution of luminous, massive galaxies ($M \sim 2 \times 10^{11} M_\odot$) at fixed stellar or dynamical mass for the first time for such a large sample size.

Despite the relatively low S/N of BOSS spectra, it is possible to measure σ for a large sample of galaxies in the range $0.2 \leq z \leq 0.7$ with a typical error $\leq 30\%$. Stellar velocity dispersions are adopted from Thomas et al. (2013).

At BOSS redshifts, effective radii are barely resolved in the SDSS imaging, and higher resolution images would be needed, which are not available for the whole sample. Therefore, we used a sub-sample of BOSS galaxies for which *HST* photometry is available as part of the COSMOS survey (Masters et al. 2011). We derived a correction to physical effective radii derived from SDSS photometry by dividing our sample in four redshift bins and searching for correlations between the ratio of the SDSS R_e and *HST*/COSMOS R_e as a function of the SDSS R_e .

We then derive dynamical mass estimates by means of a simple virial mass estimator based on galaxy effective radius and velocity dispersions within the effective radius. These total dynamical masses are compared to the total stellar masses derived by Maraston et al. (2013) studying the redshift evolution of the galaxy parameters effective radius, stellar velocity dispersion, and dynamical to stellar mass ratio M_{dyn}/M_* . We complement the SDSS-III/BOSS sample with local early-type galaxies from SDSS-II after matching their mass distributions, so that our study covers the redshift range $0.1 < z < 0.55$.

To account for the effects of the so-called progenitor bias, we compare the galaxy ages in each redshift bin and remove those galaxies from the low- z sample whose ages after evolution to the highest redshift bin would be lower than a given age threshold. As a result, we study a sample of passively evolving galaxies within a relatively narrow mass range about $M \sim 2 \times 10^{11} M_\odot$ (for a Kroupa IMF).

We find a moderate size evolution at the $\sim 1.5\sigma$ level, with galaxy radii decreasing with increasing redshift in agreement with previous results and model predictions (Oser et al. 2012; but better with Nipoti et al. 2012; Khochfar & Silk 2006 or Hopkins et al. 2009). We further observe a mild, but significant ($>2\sigma$), evolution in velocity dispersion of σ increasing with increasing redshift. The evolution of stellar velocity dispersion and effective radius together combine to an evolution of the dynamical to stellar mass ratio, such that M_{dyn}/M_* increases with decreasing redshift at $>2\sigma$ significance. We emphasize that we are probing a variation of stellar kinematics within the effective radius, hence this evolution is caused by a change of dynamical mass within $1 R_e$, even though total masses are compared.

The major sources for random and systematic errors are the size correction and the calculation of dynamical mass through the virial estimator. To assess random and systematic errors in the redshift evolution of these galaxy parameters, we perform Monte Carlo simulations perturbing the size correction, as well as the structural dependent constant of proportionality of the virial mass estimator within their errors. We present several additional tests on the impact of aperture effects (Appendix A), of unresolved multiple systems in the SDSS images (Appendix B), of a possible redshift dependence of the virial constant (Appendix C), of the progenitor bias correction (Appendix D), and of a mismatch in mass distribution between the local and the high- z galaxy samples (Appendix E). We show that, while the details and the exact strengths of the correlations between R_e , σ , and M_{dyn}/M_* with redshift vary, the general detection of a redshift evolution of these parameters is robust against the systematic uncertainties from these procedures.

Finally, we extend the present study to higher redshifts by combining our sample with high-redshift literature data (Saglia et al. 2010; van de Sande et al. 2013) so that we cover the full redshift range from $z \sim 0.1$ to $z \sim 2$. The high- z sample is consistent with the trend of decreasing M_{dyn}/M_* , even though the scatter at $z > 1$ is large. By fitting the data over the full redshift range we find $M_{\text{dyn}}/M_* \propto (1+z)^{-0.30 \pm 0.12}$. This slope is slightly shallower but well consistent with the value derived in this work from the SDSS-II and SDSS-III/BOSS data alone. Our results are clearly driven by the BOSS sample in which the large number statistics allows us to have smaller uncertainties on average quantities as a function of redshift and help to identify and quantify redshift-dependent trends. Most importantly, the evidence for a decrease of M_{dyn}/M_* with increasing redshift is reinforced further at $>2\sigma$ statistical significance, although a

larger sample of high-redshift data would be helpful to constrain the evolution M_{dyn}/M_{\star} ratio over a redshift range wider than that of BOSS.

We discuss that the increase of M_{dyn} with cosmic time is most plausibly caused by an increase of dark matter fraction within the effective radius. This evolution can be well understood through the size growth, which causes an increase of the dark matter fraction within an increasing effective radius as also predicted by galaxy formation simulations based on minor merger driven mass growth. Major mergers could also result in an evolution of M_{dyn}/M_{\star} , but of smaller amount; however, with our data we cannot constrain any difference between minor and major mergers. Finally, it is interesting to note that the steady increase of the dark matter fraction in the centers of massive galaxies with time further implies that massive galaxies in the local universe must contain some dark matter within their half-light radii, even if they are baryon dominated.

We acknowledge the anonymous referee for valuable comments that led to an improved presentation. A.B., D.T., C.M., O.S., K.M., J.P., R.T., J.J., and R.N. acknowledge STFC rolling grant ST/I001204/1 ‘‘Survey Cosmology and Astrophysics’’ for support. A.B. is indebted with Michele Cirasuolo, Lodovico Coccato, Enrico Maria Corsini, Marc Sarzi, Sadegh Khochfar, Thorsten Naab, Nicola Napolitano, Stefanie Phleps, Jens Thomas, and David Wilman for many useful discussions and suggestions.

K.L.M. acknowledges funding from The Leverhulme Trust as a 2010 Early Career Fellow.

Funding for SDSS-III has been provided by the Alfred P. Sloan Foundation, the Participating Institutions, the National Science Foundation, and the U.S. Department of Energy Office of Science. The SDSS-III Web site is <http://www.sdss3.org/>.

SDSS-III is managed by the Astrophysical Research Consortium for the Participating Institutions of the SDSS-III Collaboration including the University of Arizona, the Brazilian Participation Group, Brookhaven National Laboratory, Carnegie Mellon University, University of Florida, the French Participation Group, the German Participation Group, Harvard University, Instituto de Astrofísica de Canarias, the Michigan State/Notre Dame/JINA Participation Group, Johns Hopkins University, Lawrence Berkeley National Laboratory, Max Planck Institute for Astrophysics, Max Planck Institute for Extraterrestrial Physics, New Mexico State University, New York University, Ohio State University, Pennsylvania State University, University of Portsmouth, Princeton University, the Spanish Participation Group, University of Tokyo, University of Utah, Vanderbilt University, University of Virginia, University of Washington, and Yale University.

This research has made use of the NASA/IPAC Infrared Science Archive, which is operated by the Jet Propulsion Laboratory, California Institute of Technology, under contract with the National Aeronautics and Space Administration.

APPENDIX A

STELLAR MASSES FROM ONE EFFECTIVE RADIUS APERTURE MAGNITUDES

The dynamical mass obtained using the virial mass estimator (see Section 3.4) is based on stellar kinematics within an aperture of 1 effective radius and scaled to *total* dynamical mass via Equation (3). This quantity is compared

with the *total* stellar mass from Maraston et al. (2013) based on `cmodelmag` magnitudes. Hence, both dynamical and stellar masses are *total* masses, which ensures a consistent comparison.

Still, the total dynamical mass is derived from observations within the effective radius, while the stellar mass comes from the total stellar light. We explore therefore the possible presence of a systematic effect from the different apertures in which kinematics and stellar populations have been measured. In this test we compare M_{\star} derived from `modelmag` (rescaled to `cmodelmag`) and M_{\star} from aperture magnitudes within R_e (rescaled to `cmodelmag`), for a sub-sample of 1,000 galaxies randomly selected among the BOSS sample to assess differences between the two quantities. We derive magnitudes within the two apertures closest to the effective radius of each galaxy in the *i* band (the same R_e we used for M_{dyn} determination) and we interpolate magnitudes to derive the value we would have at R_e in each band. Aperture magnitudes are then rescaled to `cmodelmag` in the *i* band to make sure that any difference on the resulting stellar mass comes from differences in the M/L ratio and SED shape.

In detail, we collect our reference (circularized) *i*-band radii following Section 3.2.1. We estimate aperture magnitudes following the step described at http://www.sdss3.org/dr8/algorithms/magnitudes.php#photo_profile through a `casjob` query. We first derive `nprof`, from the `PhotoObjAll` table, which gives us the number of annuli (concentric circles) for which there is a measurable signal in *u*, *g*, *r*, *i*, *z* bands. From the `photoProfile` table, where azimuthally-averaged radial profiles of SDSS photo objects are listed, we create circular aperture magnitudes from `profMean` (the mean surface brightness within the annuli) and their errors (`profErr`) which are both in *nanomaggies/arcsec*². Quantities are then converted from *nanomaggie* following <http://www.sdss3.org/dr8/algorithms/magnitudes.php#nmgy>. The `bin` keyword gives the annuli from which the profile was derived (from 0 to 14) and `band` gives the selected *u*-, *g*-, *r*-, *i*-, *z*-band (from 0–4). We integrate `profMean` values within each of the two annuli close to the circularized *i*-band effective radii. In our query we impose that magnitudes are only calculated when fluxes within the aperture in *nanomaggies* are positive. We also tested whether the number of available bins for each galaxy is smaller or equal to the requested aperture. The final sample for which we are able to derive aperture magnitudes reduces to 604 galaxies due to the quality of the *u*-band data. Magnitudes in the two annuli close to R_e are then linearly interpolated to derive the value at R_e .

The histogram of the resulting difference between M_{\star} obtained using `modelmag` rescaled to `cmodelmag` and 1 R_e aperture magnitudes rescaled to `cmodelmag` is shown in Figure 12. The difference is ~ -0.08 dex in log (25%). The slightly redder populations in the more central aperture photometry (1 R_e aperture magnitude colors are redder than `modelmag` colors) leads to slightly higher M/L ratios and hence slightly higher masses. However, this result suggests that the systematic offset is small. Our M_{dyn}/M_{\star} ratios are always above 0.08 dex, so by rescaling Maraston et al. (2013) stellar masses there is no risk that M_{\star} exceeds M_{dyn} , which would be unphysical. We also note that the difference between M_{\star} obtained using `modelmag` rescaled to `cmodelmag` and 1 R_e aperture magnitudes rescaled to `cmodelmag` does not depend on redshift (the redshift variation is ~ 0.01 dex, of the order of the uncertainties on M_{\star}). Therefore, our analysis, focused on redshift variations, is not affected.

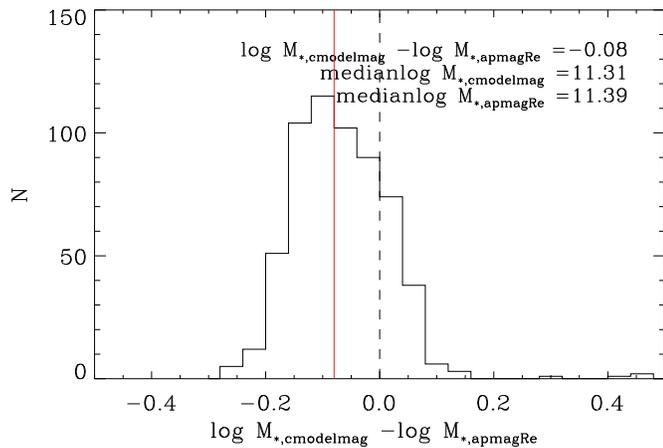


Figure 12. Comparison between stellar masses derived from `modelmag` and aperture magnitudes within $1 R_e$ both rescaled to i -band `cmodelmag` to give total mass. The median difference between the two masses is ~ -0.08 dex (25%; continuous red line). Galaxies are redder inside the effective radius leading to higher M/L ratios and hence larger total mass (when scaled to total light).

(A color version of this figure is available in the online journal.)

APPENDIX B

IMPACT OF UNRESOLVED MULTIPLE SYSTEMS IN THE SDSS IMAGES

The presence of unresolved multiple systems in the SDSS images could potentially introduce spurious trends in the measured evolution of sizes, dynamical masses and stellar masses. In this Appendix we study their influence in our size calibration and in the stellar mass estimation.

Masters et al. (2011) found that the 23% of the COSMOS/BOSS sub-sample shows two components in the COSMOS images, which are unresolved in the SDSS images. We further investigate this, considering only objects showing a size overestimation after our size calibration.

Figure 13, left panel, shows the distribution of the ratio between corrected sizes and COSMOS sizes for 206 objects part

of the COSMOS/BOSS sub-sample in the redshift range $0.2 \lesssim z \lesssim 0.7$. Those include both single objects (163 galaxies, blue filled histogram) and unresolved multiple systems (43 galaxies, red empty histogram). Our size correction allows to adjust most of the unresolved multiple systems, resulting in a distribution similar to that of the single objects but with a tail at larger ratios. A closer look at the catalog of Masters et al. (2011) (available at <http://www.icg.port.ac.uk/~mastersk/BOSSmorphologies/>) shows that when the secondary object in the system is very faint, its effect on the radius measured by the SDSS pipeline seems to be reduced, resulting in sizes not massively overestimated.

We estimate the further correction we have to apply to unresolved multiple systems, considering the 12 objects ($\sim 6\%$ of our COSMOS/BOSS sub-sample) which are in the tail, as shown by the green histogram in the middle panel of Figure 13. Although the effect of the multiple systems is likely to be redshift dependent, the small sample we are considering here does not allow a division in redshift bins. Moreover, due to the small statistic, we make use the full distribution to derive the correction factor for sizes, which ranges between a factor two and three.

Similarly to sizes, stellar masses of unresolved multiple systems could be biased too. We use the same 12 galaxies with overestimated corrected sizes to assess this effect. Under the assumption that our galaxies are passively evolving, we can estimate the bias in the stellar mass from the ratio of the fluxes of the two components resolved by the COSMOS imaging. If galaxies were not passively evolving, the luminosity would be diluted but our conclusions would not change appreciably. Fluxes were derived from the `MAG_AUTO` listed in the Zurich Structure & Morphology Catalog v1.0 available at <http://irsa.ipac.caltech.edu/data/COSMOS/datasets.html> and described in Section 3.2.2. Given that we do not have accurate information about the redshifts of the two systems we used as estimate the parent magnitudes and not absolute magnitudes. Two objects are discarded because, by visual inspection, their separation looked large enough for each of them to be considered two separated sources, suggesting a smaller bias in the SDSS magnitudes used in the stellar mass calculation

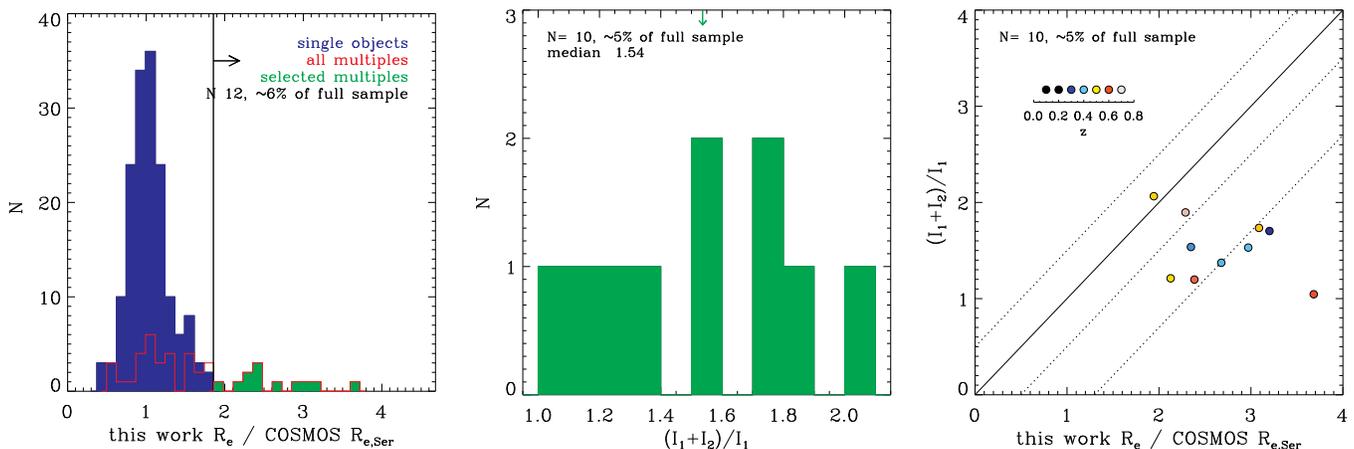


Figure 13. Left panel: distribution of the ratio between our corrected SDSS R_e and COSMOS $R_{e,Ser}$ (blue filled histogram) and distribution of unresolved multiple systems (red empty histogram). The green filled histogram shows the sample of unresolved multiple systems which were used to study the impact of these objects on our results. Middle panel: distribution of the ratio between the fluxes (masses) of the two components which are resolved in the COSMOS photometry. Right panel: comparison between the size excess and the flux (mass) excess in our sample of 10 multiple systems coded as a function of redshift. By using sizes from SDSS we tend to overestimate sizes of a factor 2–3 and fluxes of a factor ~ 1.54 . The continuous line corresponds to the one-to-one relation and dotted lines show the intercepts ± 0.5 and $1.3x$.

(A color version of this figure is available in the online journal.)

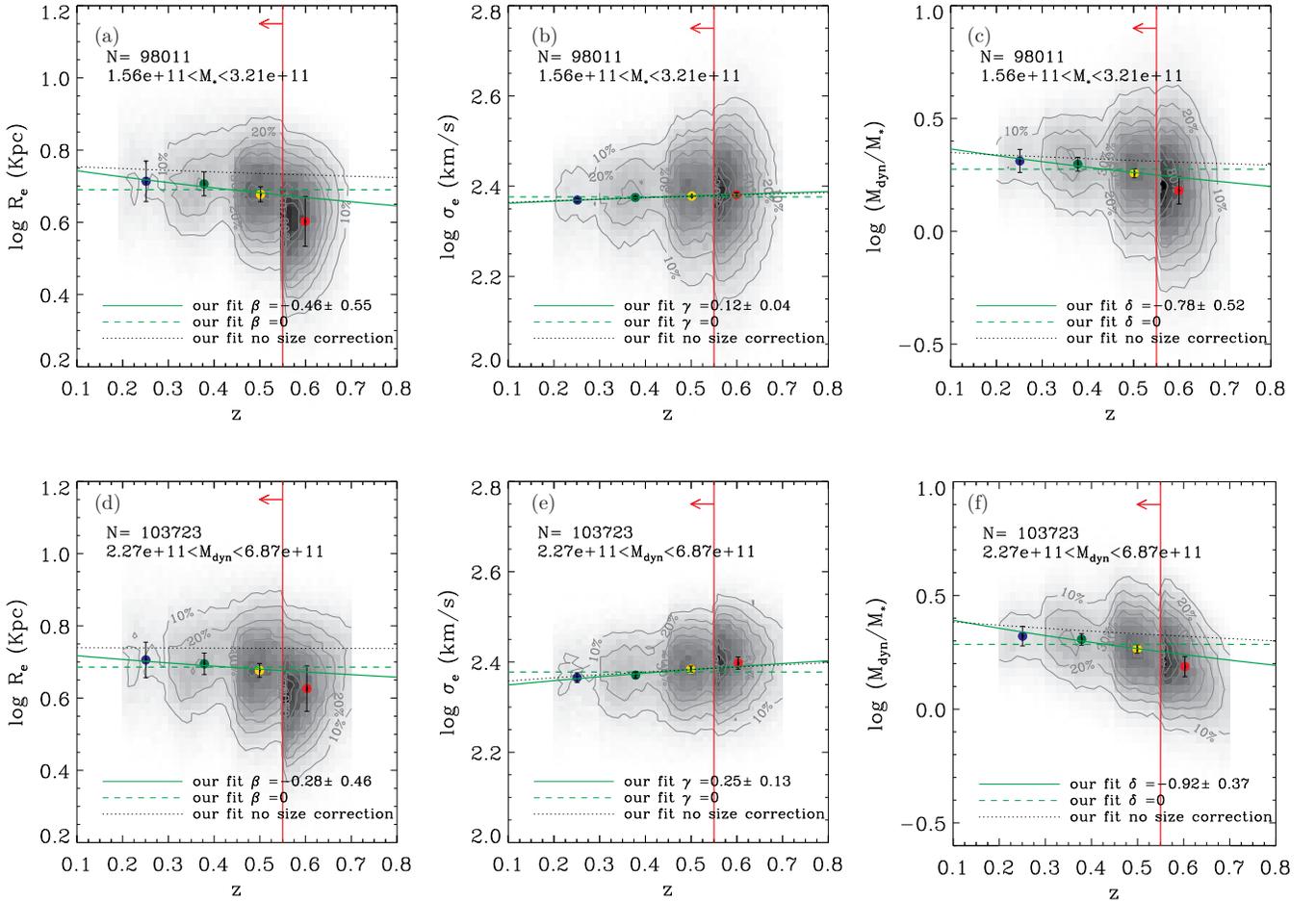


Figure 14. Same as Figure 8, but for the SDSS-III/BOSS sample only and with a redshift-dependent β_{dyn} (see Equation (3)).

(A color version of this figure is available in the online journal.)

Table 3
Fitting Parameters for the BOSS Sample between $0.2 < z < 0.55$, in which the Progenitor-bias Correction is Applied

Parameter	M_*				M_{dyn}			
	Constant β_{dyn}		z -variable β_{dyn}		Constant β_{dyn}		z -variable β_{dyn}	
	Slope	Zero Point	Slope	Zero Point	Slope	Zero Point	Slope	Zero Point
R_e	-0.46 ± 0.56	0.76 ± 0.09	-0.46 ± 0.55	0.76 ± 0.09	-0.32 ± 0.47	0.73 ± 0.08	-0.28 ± 0.46	0.73 ± 0.08
σ_e	0.12 ± 0.04	2.36 ± 0.006	0.12 ± 0.04	2.36 ± 0.006	0.20 ± 0.13	2.35 ± 0.02	0.25 ± 0.13	2.34 ± 0.02
M_{dyn}/M_*	-0.53 ± 0.48	0.36 ± 0.08	-0.78 ± 0.52	0.40 ± 0.08	-0.77 ± 0.35	0.41 ± 0.06	-0.92 ± 0.37	0.43 ± 0.06

Notes. Uncertainties on each parameter are 1σ errors derived from Monte Carlo simulations. The relation we fitted for R_e is $\log R_e = \log R_{e,0} + \beta(1+z)$, for σ_e is $\log \sigma_e = \log \sigma_{e,0} + \gamma(1+z)$, and for M_{dyn}/M_* is $\log(M_{\text{dyn}}/M_*) = \log(M_{\text{dyn}}/M_*)_0 + \delta(1+z)$.

(separation $> 2 \times \text{FWHM}_{\text{seeing}}$). This retains 10 objects ($\sim 5\%$ of the total sample of COSMOS/BOSS galaxies). The distribution of the relative flux ratio of the objects which are resolved in the COSMOS images, $(I_1 + I_2)/I_1$, is shown in the middle panel of Figure 13, where I_1 is the flux of the primary brightest object in COSMOS, and I_2 is the flux of the secondary objects within 2 times the typical FWHM of the SDSS seeing ($\sim 1''.1$; see Section 3.2.1 for details). The average correction factor for fluxes (or masses) is about 1.54.

Although stellar velocity dispersions, to some extent, can be biased too, the estimation of this effect is not trivial and cannot be done without additional spectroscopy, therefore we

disregard any possible biases to stellar velocity dispersions due to unresolved multiple in our analysis.

The right panel of Figure 13 compares the flux (i.e., mass) correction with the size correction for final sample of 10 unresolved multiple systems used here. Typical size corrections range between a factor two and three, whereas the flux (mass) correction is much smaller (~ 1.54) as shown by the dotted lines parallel to the continuous one-to-one relation in a manner which does not seem to change with redshift (but the statistic is too small to make strong statements).

We finally apply the redshift-independent size and mass corrections to a randomly selected 6% sub-sample of BOSS

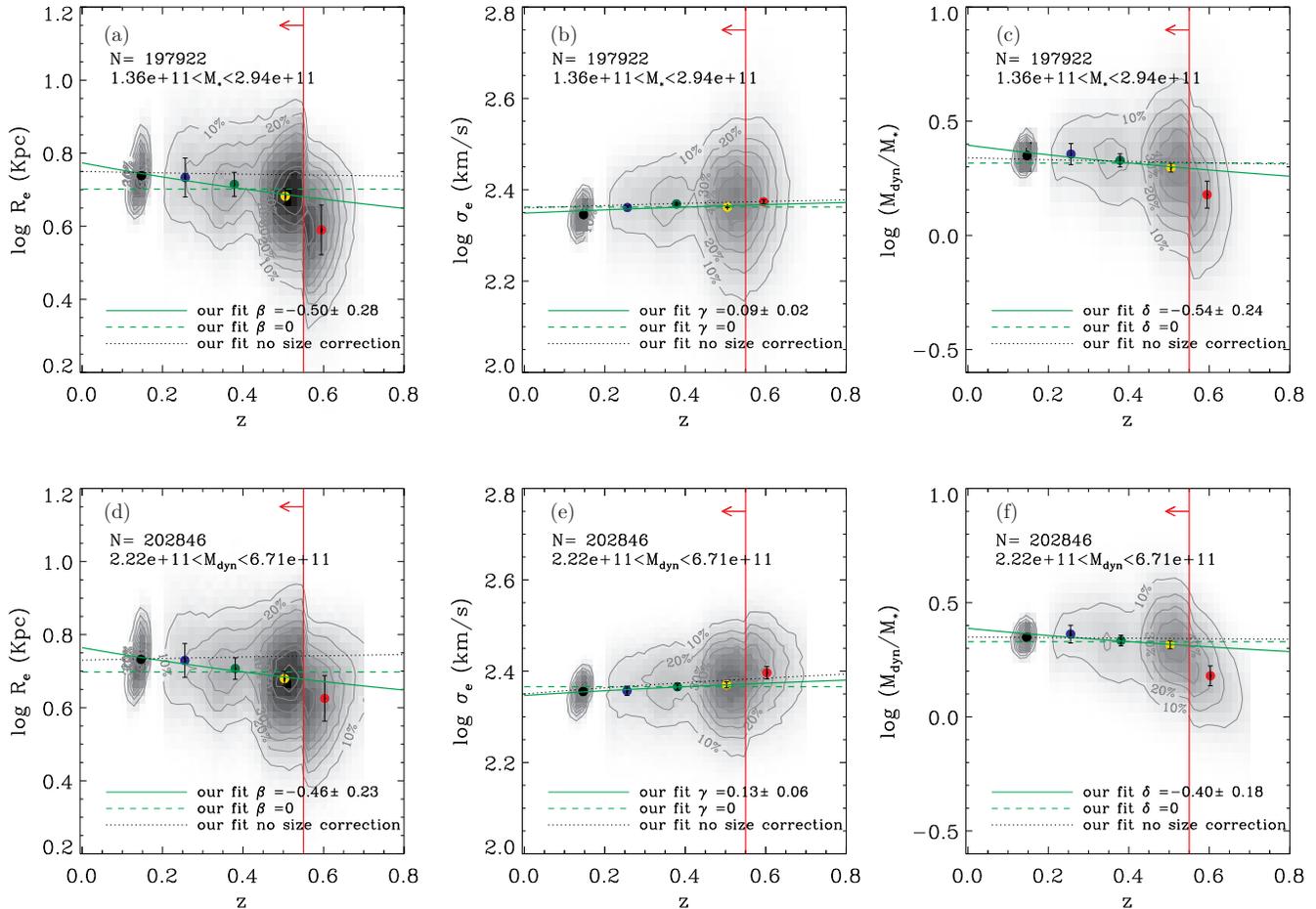


Figure 15. Same as Figure 8, but no correction for progenitor bias.
(A color version of this figure is available in the online journal.)

galaxies and analyze the effects on our results. The corrections were also included in our Monte Carlo estimation of systematic errors. The resulted combined effect on sizes and masses is negligible and our results do not change substantially considering this additional correction or not (if not improve them of about $<2\%$). Figures 8, 9, 10, and 11 and Table 2 include this additional correction.

APPENDIX C

TESTS ON EVOLUTION: BOSS SAMPLE AND USE OF CONSTANT OR REDSHIFT-DEPENDENT β_{dyn}

In this Appendix, we show that our results do not change when we analyze scaling relations only using the BOSS sample. Table 3 (case in which β_{dyn} is assumed constant with redshift) shows that the trends we find in Section 4 are maintained, although the significance of the results is reduced. Errors from the Monte Carlo simulations for the BOSS sample are 50%–60% larger than using the combined sample of local and BOSS data (see Figure 8 and Table 2). The significance is particularly reduced for the size evolution. We conclude that the addition of the local early-type sample does not affect our results but helps constraining the fits. Results are summarized in Table 3.

Sérsic indices from *HST* photometry are available as part of the BOSS/COSMOS sub-sample (see Section 3.4.1).

This allows us to assess the effect of a redshift-dependent $\beta_{\text{dyn}}(n)$ parameter through our Monte Carlo simulations. The results of this test are shown in Figure 14 and summarized in

Table 3 (columns with z -variable β_{dyn}). By using a redshift-dependent β_{dyn} on Equation (3) for M_{dyn} , the evolution of the M_{dyn}/M_* ratio becomes more significant than using a constant β_{dyn} (more conservative results). The difference on the slopes, however, is very small ($\sim 10\%$) suggesting that $\beta_{\text{dyn}}(n)$ does not strongly evolve in the BOSS redshift range and any possible variation would not affect our results. Table 3 summarizes the comparison between a redshift dependent or a constant β_{dyn} .

The analysis above can be done only for the COSMOS/BOSS sample, because structural parameters are measured in a similar way and on the same images.

We use the photometric catalog of Simard et al. (2011) for SDSS local galaxies to compare structural parameters (i.e., Sérsic indices) of the local and BOSS/COSMOS sample. Simard et al. (2011) presented new structural parameters from both single Sérsic fits and bulge-to-disk decomposition on the full DR7 sample in the g and r bands. For our tests we chose data from single Sérsic fit to be consistent with COSMOS (in particular around the transition redshift $z \sim 0.2$).

The Simard et al. (2011) and COSMOS catalogs have a limited number of objects in common. For $n < 4$ we found a fair agreement between structural parameters in the two catalog, whereas for $n > 4$, Simard et al. (2011) Sérsic indices are larger than COSMOS ones. By using median Simard et al. (2011) Sérsic indices for our early-type sample, we reproduce the commonly used $\beta_{\text{dyn}}(n) \sim 5$ for local galaxies. Sérsic indices of the two samples probably differ because they were derived with

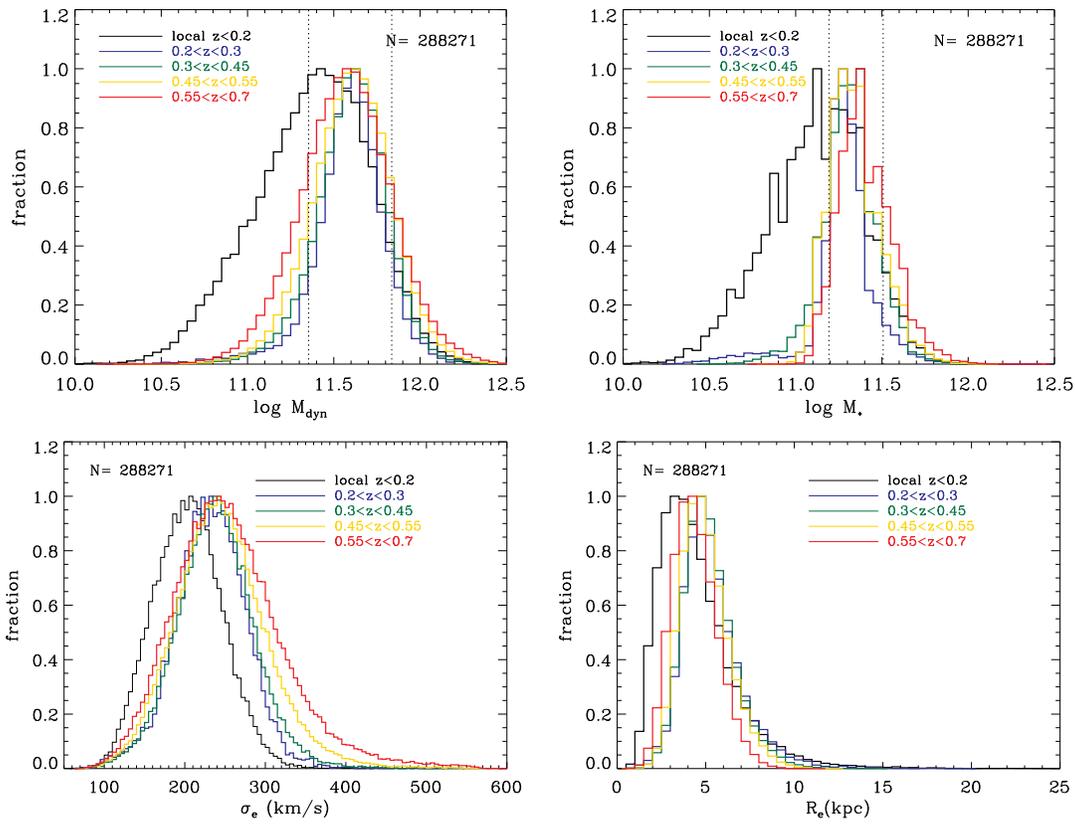


Figure 16. Same as Figure 7, but no homogenization between the mass distributions of the local sample from SDSS-II and the sample from SDSS-III/BOSS. (A color version of this figure is available in the online journal.)

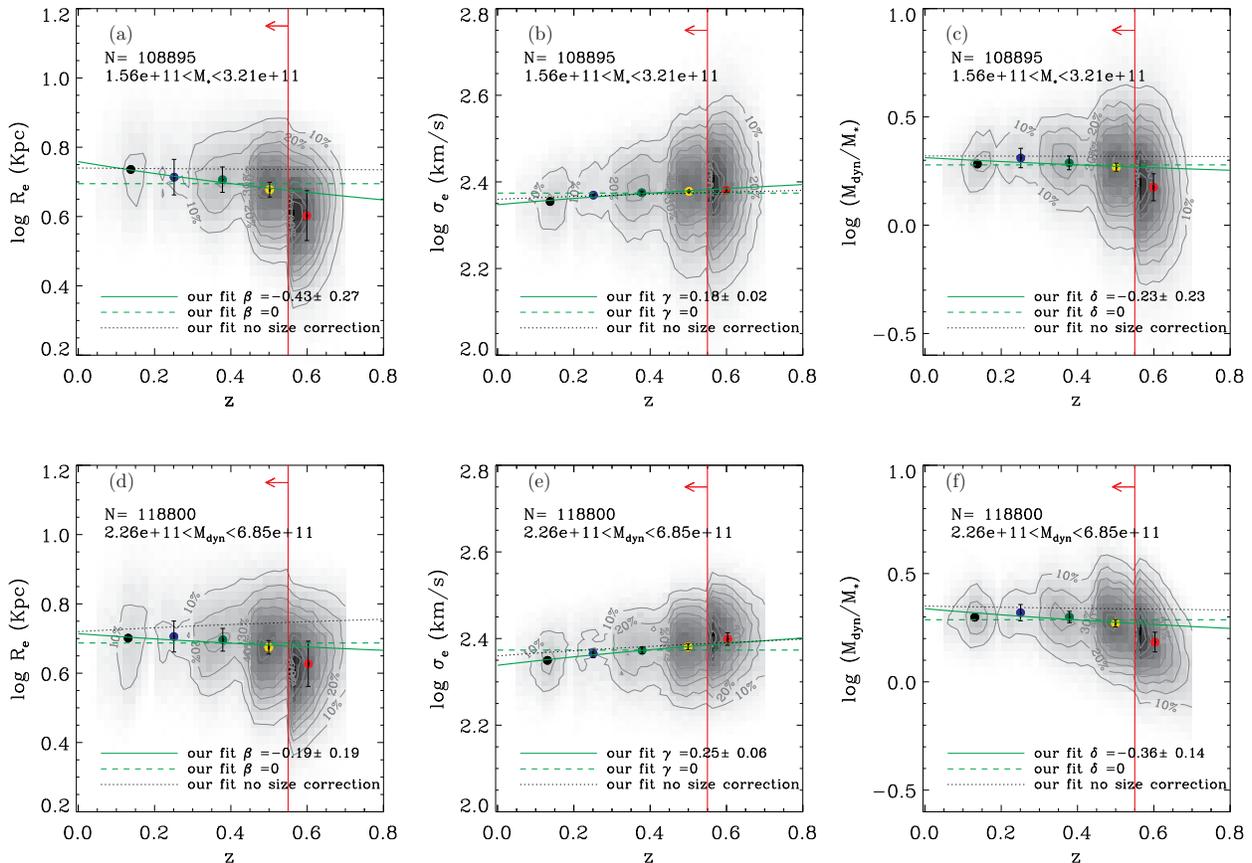


Figure 17. Same as Figure 8, but no homogenization between the mass distributions of the local sample from SDSS-II and the sample from SDSS-III/BOSS. (A color version of this figure is available in the online journal.)

Table 4Fitting Parameters for the Redshift Evolution of Galaxy Parameters between $0.1 \leq z \leq 0.55$ without Progenitor Bias Correction

Parameter	M_*		M_{dyn}	
	Slope	Zero Point	Slope	Zero Point
R_e	-0.50 ± 0.28	0.78 ± 0.04	-0.46 ± 0.23	0.76 ± 0.03
σ_e	0.09 ± 0.02	2.36 ± 0.00	0.13 ± 0.06	2.35 ± 0.01
M_{dyn}/M_*	-0.54 ± 0.24	0.40 ± 0.03	-0.40 ± 0.18	0.38 ± 0.02

Notes. Uncertainties on each parameter are 1σ errors derived from Monte Carlo Simulations. The relation we fitted for R_e is $\log R_e = \log R_{e,0} + \beta(1+z)$, for σ_e is $\log \sigma_e = \log \sigma_{e,0} + \gamma(1+z)$, and for M_{dyn}/M_* is $\log(M_{\text{dyn}}/M_*) = \log(M_{\text{dyn}}/M_*)_0 + \delta(1+z)$.

different images and approaches and there is no clear correction to be applied to this.

For this reason, in our analysis, we do not combine information about the Sérsic indices from Simard et al. (2011) and COSMOS. Instead, we adopt the redshift-independent β_{dyn} of the BOSS galaxies also for the SDSS sample.

An explanation of these differences is beyond the scope of this paper.

APPENDIX D

TESTS ON EVOLUTION: LOCAL EARLY-TYPES & BOSS SAMPLES WITHOUT PROGENITOR BIAS CORRECTION

In Section 3.6, we discuss the correction for progenitor bias applied in this work. To study the impact of this correction on our analysis, we have performed a re-analysis for a sample without progenitor bias correction. Figure 15 shows the redshift evolution of the galaxy parameters size, stellar velocity dispersion and M_{dyn}/M_* without progenitor bias correction. The fit parameters are given in Table 4. Comparing to Figure 8 and Table 2 in the main text, it can be seen that generally the results are fairly stable against the progenitor-bias correction, significant evolution with redshift is still detected for all three parameters. We infer that the main conclusions of this paper do not critically depend on the progenitor bias correction.

APPENDIX E

TESTS ON EVOLUTION: STELLAR MASS DISTRIBUTION OF LOCAL EARLY-TYPES & BOSS GALAXIES

As described in Section 3.5, we homogenize the stellar mass distributions between the local sample from SDSS-II and the high- z sample from SDSS-III/BOSS by selecting a local subsample that matches the mass distribution of the BOSS sample. In this Appendix, we present a re-analysis in which we do not apply this homogenization. The distributions of stellar masses, dynamical masses, stellar velocity dispersions, and effective radii are shown in Figure 16. Figure 17 shows the redshift evolution of the galaxy parameters size, stellar velocity dispersion, and M_{dyn}/M_* ; the fit parameters are summarized in Table 5. It can be seen that our finding of redshift evolution remains intact for all three parameters, and the significance of the slopes only changes slightly.

REFERENCES

Abazajian, K. N., Adelman-McCarthy, J. K., Agüeros, M. A., et al. 2009, *ApJS*, **182**, 543

Table 5Fitting Parameters for the Redshift Evolution of Galaxy Parameters between $0.1 \leq z \leq 0.55$ without Matching Mass Distributions between the Local SDSS-II and the SDSS-III/BOSS Samples

Parameter	M_*		M_{dyn}	
	Slope	Zero Point	Slope	Zero Point
R_e	-0.43 ± 0.27	0.76 ± 0.04	-0.19 ± 0.19	0.72 ± 0.03
σ_e	0.18 ± 0.02	2.35 ± 0.00	0.25 ± 0.06	2.34 ± 0.01
M_{dyn}/M_*	-0.23 ± 0.23	0.32 ± 0.03	-0.36 ± 0.14	0.34 ± 0.02

Notes. Uncertainties on each parameter are 1σ errors derived from Monte Carlo simulations. The relation we fitted for R_e is $\log R_e = \log R_{e,0} + \beta(1+z)$, for σ_e is $\log \sigma_e = \log \sigma_{e,0} + \gamma(1+z)$, and for M_{dyn}/M_* is $\log(M_{\text{dyn}}/M_*) = \log(M_{\text{dyn}}/M_*)_0 + \delta(1+z)$.

- Ahn, C. P., Alexandroff, R., Allende Prieto, C., et al. 2012, *ApJS*, **203**, 21
Aihara, H., Allende Prieto, C., An, D., et al. 2011, *ApJS*, **193**, 29
Anderson, L., Aubourg, E., Bailey, S., et al. 2012, *MNRAS*, **427**, 3435
Auger, M. W., Treu, T., Bolton, A. S., et al. 2010a, *ApJ*, **724**, 511
Auger, M. W., Treu, T., Gavazzi, R., et al. 2010b, *ApJL*, **721**, L163
Beifiori, A., Courteau, S., Corsini, E. M., & Zhu, Y. 2012, *MNRAS*, **419**, 2497
Beifiori, A., Maraston, C., Thomas, D., & Johansson, J. 2011, *A&A*, **531**, A109
Bell, E. F., & de Jong, R. S. 2001, *ApJ*, **550**, 212
Belli, S., Newman, A. B., & Ellis, R. S. 2014, *ApJ*, **783**, 117
Bender, R. 1990, *A&A*, **229**, 441
Bender, R., Burstein, D., & Faber, S. M. 1992, *ApJ*, **399**, 462
Bender, R., Burstein, D., & Faber, S. M. 1993, *ApJ*, **411**, 153
Bender, R., Saglia, R. P., & Gerhard, O. E. 1994, *MNRAS*, **269**, 785
Bender, R., Saglia, R. P., Ziegler, B., et al. 1998, *ApJ*, **493**, 529
Bernardi, M., Shankar, F., Hyde, J. B., et al. 2010, *MNRAS*, **404**, 2087
Bernardi, M. 2009, *MNRAS*, **395**, 1491
Bernardi, M., Hyde, J. B., Fritz, A., et al. 2008, *MNRAS*, **391**, 1191
Bernardi, M., Hyde, J. B., Sheth, R. K., Miller, C. J., & Nichol, R. C. 2007, *AJ*, **133**, 1741
Bernardi, M., Sheth, R. K., Annis, J., et al. 2003, *AJ*, **125**, 1817
Bernardi, M., Sheth, R. K., Nichol, R. C., et al. 2006, *AJ*, **131**, 2018
Bertin, G., Ciotti, L., & Del Principe, M. 2002, *A&A*, **386**, 149
Bezanson, R., van Dokkum, P. G., Tal, T., et al. 2009, *ApJ*, **697**, 1290
Bezanson, R., van Dokkum, P., van de Sande, J., Franx, M., & Kriek, M. 2013b, *ApJL*, **764**, L8
Bezanson, R., van Dokkum, P. G., van de Sande, J., et al. 2013a, *ApJL*, **779**, L21
Blanton, M. R., Kazin, E., Muna, D., Weaver, B. A., & Price-Whelan, A. 2011, *AJ*, **142**, 31
Bolton, A. S., Brownstein, J. R., Kochanek, C. S., et al. 2012a, *ApJ*, **757**, 82
Bolton, A. S., Burles, S., Koopmans, L. V. E., Treu, T., & Moustakas, L. A. 2006, *ApJ*, **638**, 703
Bolton, A. S., Burles, S., Koopmans, L. V. E., et al. 2008, *ApJ*, **682**, 964
Bolton, A. S., Schlegel, D. J., Aubourg, É., et al. 2012b, *AJ*, **144**, 144
Bolzonella, M., Kovač, K., Pozzetti, L., et al. 2010, *A&A*, **524**, A76
Boylan-Kolchin, M., Ma, C.-P., & Quataert, E. 2005, *MNRAS*, **362**, 184
Brownstein, J. R., Bolton, A. S., Schlegel, D. J., et al. 2012, *ApJ*, **744**, 41
Bruzual, G., & Charlot, S. 2003, *MNRAS*, **344**, 1000
Buitrago, F., Trujillo, I., Conselice, C. J., et al. 2008, *ApJL*, **687**, L61
Caon, N., Capaccioli, M., & D'Onofrio, M. 1993, *MNRAS*, **265**, 1013
Cappellari, M., Bacon, R., Bureau, M., et al. 2006, *MNRAS*, **366**, 1126
Cappellari, M., di Serego Alighieri, S., Cimatti, A., et al. 2009, *ApJL*, **704**, L34
Cappellari, M., & Emsellem, E. 2004, *PASP*, **116**, 138
Cappellari, M., McDermid, R. M., Alatalo, K., et al. 2012, *Natur*, **484**, 485
Cappellari, M., McDermid, R. M., Alatalo, K., et al. 2013a, *MNRAS*, **432**, 1862
Cappellari, M., Scott, N., Alatalo, K., et al. 2013b, *MNRAS*, **432**, 1709
Carrasco, E. R., Conselice, C. J., & Trujillo, I. 2010, *MNRAS*, **405**, 2253
Cenarro, A. J., & Trujillo, I. 2009, *ApJL*, **696**, L43
Chabrier, G. 2003, *ApJL*, **586**, L133
Chen, Y.-M., Kauffmann, G., Tremonti, C. A., et al. 2012, *MNRAS*, **421**, 314
Cimatti, A., Cassata, P., Pozzetti, L., et al. 2008, *A&A*, **482**, 21
Cimatti, A., Nipoti, C., & Cassata, P. 2012, *MNRAS*, **422**, L62
Ciotti, L., Lanzoni, B., & Renzini, A. 1996, *MNRAS*, **282**, 1
Conroy, C., Dutton, A. A., Graves, G. J., Mendel, J. T., & van Dokkum, P. G. 2013, *ApJL*, **776**, L26

- Conroy, C., & van Dokkum, P. G. 2012, *ApJ*, **760**, 71
- Daddi, E., Renzini, A., Pirzkal, N., et al. 2005, *ApJ*, **626**, 680
- Damjanov, I., Chilingarian, I., Hwang, H. S., & Geller, M. J. 2013, *ApJL*, **775**, L48
- Damjanov, I., McCarthy, P. J., Abraham, R. G., et al. 2009, *ApJ*, **695**, 101
- Dawson, K. S., Schlegel, D. J., Ahn, C. P., et al. 2013, *AJ*, **145**, 10
- de Zeeuw, P. T., Bureau, M., Emsellem, E., et al. 2002, *MNRAS*, **329**, 513
- Djorgovski, S., & Davis, M. 1987, *ApJ*, **313**, 59
- Dressler, A., Lynden-Bell, D., Burstein, D., et al. 1987, *ApJ*, **313**, 42
- Dutton, A. A., Conroy, C., van den Bosch, F. C., et al. 2011, *MNRAS*, **416**, 322
- Dutton, A. A., Macciò, A. V., Mendel, J. T., & Simard, L. 2013, *MNRAS*, **432**, 2496
- Dutton, A. A., Mendel, J. T., & Simard, L. 2012, *MNRAS*, **422**, L33
- Eisenstein, D. J., Annis, J., Gunn, J. E., et al. 2001, *AJ*, **122**, 2267
- Eisenstein, D. J., Weinberg, D. H., Agol, E., et al. 2011, *AJ*, **142**, 72
- Franx, M., van Dokkum, P. G., Schreiber, N. M. F., et al. 2008, *ApJ*, **688**, 770
- Fukugita, M., Ichikawa, T., Gunn, J. E., et al. 1996, *AJ*, **111**, 1748
- Gerhard, O., Kronawitter, A., Saglia, R. P., & Bender, R. 2001, *AJ*, **121**, 1936
- Grillo, C. 2010, *ApJ*, **722**, 779
- Gunn, J. E., Carr, M., Rockosi, C., et al. 1998, *AJ*, **116**, 3040
- Gunn, J. E., Siegmund, W. A., Mannery, E. J., et al. 2006, *AJ*, **131**, 2332
- Guo, Y., McIntosh, D. H., Mo, H. J., et al. 2009, *MNRAS*, **398**, 1129
- Hilz, M., Naab, T., & Ostriker, J. P. 2013, *MNRAS*, **429**, 2924
- Hilz, M., Naab, T., Ostriker, J. P., et al. 2012, *MNRAS*, **425**, 3119
- Hogg, D. W. 1999, arXiv:astro-ph/9905116
- Hopkins, P. F., Hernquist, L., Cox, T. J., Keres, D., & Wuyts, S. 2009, *ApJ*, **691**, 1424
- Hopkins, P. F., Murray, N., & Thompson, T. A. 2009, *MNRAS*, **398**, 303
- Houghton, R. C. W., Davies, R. L., Dalla Bontà, E., & Masters, R. 2012, *MNRAS*, **423**, 256
- Hyde, J. B., & Bernardi, M. 2009a, *MNRAS*, **394**, 1978
- Hyde, J. B., & Bernardi, M. 2009b, *MNRAS*, **396**, 1171
- Ilbert, O., Salvato, M., Le Floc'h, E., et al. 2010, *ApJ*, **709**, 644
- Johansson, P. H., Naab, T., & Ostriker, J. P. 2012, *ApJ*, **754**, 115
- Jørgensen, I., Chiboucas, K., Flint, K., et al. 2006, *ApJL*, **639**, L9
- Jorgensen, I., Franx, M., & Kjaergaard, P. 1995, *MNRAS*, **276**, 1341
- Khochfar, S., & Silk, J. 2006, *ApJL*, **648**, L21
- Koekemoer, A. M., Aussel, H., Calzetti, D., et al. 2007, *ApJS*, **172**, 196
- Kroupa, P. 2001, *MNRAS*, **322**, 231
- Lauer, T. R., Faber, S. M., Richstone, D., et al. 2007, *ApJ*, **662**, 808
- Leauthaud, A., Massey, R., Kneib, J.-P., et al. 2007, *ApJS*, **172**, 219
- Longhetti, M., Saracco, P., Severgnini, P., et al. 2007, *MNRAS*, **374**, 614
- López-Sanjuan, C., Le Fèvre, O., Ilbert, O., et al. 2012, *A&A*, **548**, A7
- Mancini, C., Daddi, E., Renzini, A., et al. 2010, *MNRAS*, **401**, 933
- Mancini, C., Matute, I., Cimatti, A., et al. 2009, *A&A*, **500**, 705
- Mandelbaum, R., Hirata, C. M., Leauthaud, A., Massey, R. J., & Rhodes, J. 2012, *MNRAS*, **420**, 1518
- Maraston, C. 2005, *MNRAS*, **362**, 799
- Maraston, C., Daddi, E., Renzini, A., et al. 2006, *ApJ*, **652**, 85
- Maraston, C., Pforr, J., Henriques, B. M., et al. 2013, *MNRAS*, **435**, 2764
- Maraston, C., Pforr, J., Renzini, A., et al. 2010, *MNRAS*, **407**, 830
- Maraston, C., & Strömbäck, G. 2011, *MNRAS*, **418**, 2785
- Maraston, C., Strömbäck, G., Thomas, D., Wake, D. A., & Nichol, R. C. 2009, *MNRAS*, **394**, L107
- Markwardt, C. B. 2009, in ASP Conf. Ser. 411, *Astronomical Data Analysis Software and Systems XVIII*, ed. D. A. Bohlender, D. Durand, & P. Dowler (San Francisco, CA: ASP), 251
- Masters, K. L., Maraston, C., Nichol, R. C., et al. 2011, *MNRAS*, **418**, 1055
- Mehlert, D., Thomas, D., Saglia, R. P., Bender, R., & Wegner, G. 2003, *A&A*, **407**, 423
- Naab, T., Johansson, P. H., & Ostriker, J. P. 2009, *ApJL*, **699**, L178
- Naab, T., Johansson, P. H., Ostriker, J. P., & Efstathiou, G. 2007, *ApJ*, **658**, 710
- Napolitano, N. R., Romanowsky, A. J., & Tortora, C. 2010, *MNRAS*, **405**, 2351
- Newman, A. B., Ellis, R. S., Bundy, K., & Treu, T. 2012, *ApJ*, **746**, 162
- Newman, A. B., Ellis, R. S., Treu, T., & Bundy, K. 2010, *ApJL*, **717**, L103
- Nipoti, C., Treu, T., Leauthaud, A., et al. 2012, *MNRAS*, **422**, 1714
- Onodera, M., Daddi, E., Gobat, R., et al. 2010, *ApJL*, **715**, L6
- Onodera, M., Renzini, A., Carollo, M., et al. 2012, *ApJ*, **755**, 26
- Oser, L., Naab, T., Ostriker, J. P., & Johansson, P. H. 2012, *ApJ*, **744**, 63
- Padmanabhan, N., Seljak, U., Strauss, M. A., et al. 2004, *NA*, **9**, 329
- Pforr, J., Maraston, C., & Tonini, C. 2012, *MNRAS*, **422**, 3285
- Poggianti, B. M., Calvi, R., Bindoni, D., et al. 2013, *ApJ*, **762**, 77
- Renzini, A. 2006, *ARA&A*, **44**, 141
- Renzini, A., & Ciotti, L. 1993, *ApJL*, **416**, L49
- Rodríguez-Merino, L. H., Chavez, M., Bertone, E., & Buzzoni, A. 2005, *ApJ*, **626**, 411
- Ross, A. J., Ho, S., Cuesta, A. J., et al. 2011, *MNRAS*, **417**, 1350
- Saglia, R. P., Bertschinger, E., Baggeley, G., et al. 1993, *MNRAS*, **264**, 961
- Saglia, R. P., Bertschinger, E., Baggeley, G., et al. 1997, *ApJS*, **109**, 79
- Saglia, R. P., Sánchez-Blázquez, P., Bender, R., et al. 2010, *A&A*, **524**, A6
- Salpeter, E. E. 1955, *ApJ*, **121**, 161
- Sánchez-Blázquez, P., Peletier, R. F., Jiménez-Vicente, J., et al. 2006, *MNRAS*, **371**, 703
- Saracco, P., Casati, A., Gargiulo, A., et al. 2014, *A&A*, in press (arXiv:1401.5600)
- Saracco, P., Longhetti, M., & Andreon, S. 2009, *MNRAS*, **392**, 718
- Sargent, M. T., Carollo, C. M., Lilly, S. J., et al. 2007, *ApJS*, **172**, 434
- Sarzi, M., Falcón-Barroso, J., Davies, R. L., et al. 2006, *MNRAS*, **366**, 1151
- Scarlata, C., Carollo, C. M., Lilly, S., et al. 2007, *ApJS*, **172**, 406
- Schulz, A. E., Mandelbaum, R., & Padmanabhan, N. 2010, *MNRAS*, **408**, 1463
- Shen, S., Mo, H. J., White, S. D. M., et al. 2003, *MNRAS*, **343**, 978
- Shetty, S., & Cappellari, M. 2014, *ApJL*, **786**, L10
- Simard, L. 1998, in ASP Conf. Ser. 145, *Astronomical Data Analysis Software and Systems VII*, ed. R. Albrecht, R. N. Hook, & H. A. Bushouse (San Francisco, CA: ASP), 108
- Simard, L., Mendel, J. T., Patton, D. R., Ellison, S. L., & McConnachie, A. W. 2011, *ApJS*, **196**, 11
- Smee, S. A., Gunn, J. E., Uomoto, A., et al. 2013, *AJ*, **146**, 32
- Stoughton, C., Lupton, R. H., Bernardi, M., et al. 2002, *AJ*, **123**, 485
- Szomoru, D., Franx, M., & van Dokkum, P. G. 2012, *ApJ*, **749**, 121
- Taylor, E. N., Franx, M., Brinchmann, J., van der Wel, A., & van Dokkum, P. G. 2010, *ApJ*, **722**, 1
- Thomas, D., Steele, O., Maraston, C., et al. 2013, *MNRAS*, **431**, 1383
- Thomas, J., Saglia, R. P., Bender, R., et al. 2011, *MNRAS*, **415**, 545
- Tiret, O., Salucci, P., Bernardi, M., Maraston, C., & Pforr, J. 2011, *MNRAS*, **411**, 1435
- Toft, S., Gallazzi, A., Zirm, A., et al. 2012, *ApJ*, **754**, 3
- Toft, S., van Dokkum, P., Franx, M., et al. 2007, *ApJ*, **671**, 285
- Tojeiro, R., Percival, W. J., Wake, D. A., et al. 2012, *MNRAS*, **424**, 136
- Treu, T., Auger, M. W., Koopmans, L. V. E., et al. 2010, *ApJ*, **709**, 1195
- Treu, T., Ellis, R. S., Liao, T. X., et al. 2005, *ApJ*, **633**, 174
- Trujillo, I., Conselice, C. J., Bundy, K., et al. 2007, *MNRAS*, **382**, 109
- Trujillo, I., Erwin, P., Asensio Ramos, A., & Graham, A. W. 2004, *AJ*, **127**, 1917
- Trujillo, I., Ferreras, I., & de La Rosa, I. G. 2011, *MNRAS*, **415**, 3903
- Trujillo, I., Feulner, G., Goranova, Y., et al. 2006a, *MNRAS*, **373**, L36
- Trujillo, I., Förster Schreiber, N. M., Rudnick, G., et al. 2006b, *ApJ*, **650**, 18
- Valentinuzzi, T., Poggianti, B. M., Saglia, R. P., et al. 2010a, *ApJL*, **721**, L19
- Valentinuzzi, T., Fritz, J., Poggianti, B. M., et al. 2010b, *ApJ*, **712**, 226
- van de Sande, J., Kriek, M., Franx, M., et al. 2011, *ApJL*, **736**, L9
- van de Sande, J., Kriek, M., Franx, M., et al. 2013, *ApJ*, **771**, 85
- van der Marel, R. P., & van Dokkum, P. G. 2007, *ApJ*, **668**, 756
- van der Wel, A., Franx, M., van Dokkum, P. G., et al. 2005, *ApJ*, **631**, 145
- van der Wel, A., Holden, B. P., Zirm, A. W., et al. 2008, *ApJ*, **688**, 48
- van Dokkum, P. G., & Conroy, C. 2012, *ApJ*, **760**, 70
- van Dokkum, P. G., & Franx, M. 2001, *ApJ*, **553**, 90
- van Dokkum, P. G., Franx, M., Kelson, D. D., & Illingworth, G. D. 1998, *ApJL*, **504**, L17
- van Dokkum, P. G., Franx, M., Kriek, M., et al. 2008, *ApJL*, **677**, L5
- van Dokkum, P. G., Kriek, M., & Franx, M. 2009, *Natur*, **460**, 717
- Wegner, G. A., Corsini, E. M., Thomas, J., et al. 2012, *AJ*, **144**, 78
- York, D. G., Adelman, J., Anderson, J. E., Jr., et al. 2000, *AJ*, **120**, 1579
- Zaritsky, D., Zabludoff, A. I., & Gonzalez, A. H. 2008, *ApJ*, **682**, 68
- Zirm, A. W., van der Wel, A., Franx, M., et al. 2007, *ApJ*, **656**, 66